

THE SPLASH SURVEY: A SPECTROSCOPIC PORTRAIT OF ANDROMEDA’S GIANT SOUTHERN STREAM

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ABSTRACT

The giant southern stream (GSS) is the most prominent tidal debris feature in M31’s stellar halo and covers a significant fraction of its southern quadrant. The GSS is a complex structure composed of a relatively metal-rich, high-surface-brightness “core” and a lower metallicity, lower-surface-brightness “envelope.” We present spectroscopy of red giant stars in six fields in the vicinity of M31’s GSS (including four new fields and improved spectroscopic reductions for two previously published fields) and one field on stream C, an arc-like feature seen in star-count maps on M31’s southeast minor axis at $R \sim 60$ kpc. These data are part of our ongoing Spectroscopic and Photometric Landscape of Andromeda’s Stellar Halo survey of M31 using the DEIMOS instrument on the Keck II 10 m telescope. Several GSS-related findings and measurements are presented here. We present the innermost kinematical detection of the GSS core to date ($R = 17$ kpc). This field also contains the inner continuation of a second kinematically cold component that was originally seen in a GSS core field at $R \sim 21$ kpc. The velocity gradients of the GSS and the second component in the combined data set are parallel over a range of $\Delta R = 7$ kpc, suggesting that this may represent a bifurcation in the line-of-sight velocities of GSS stars. We present the first kinematical detection of substructure in the GSS envelope (S quadrant, $R \sim 58$ kpc). Using kinematically identified samples, we show that the envelope debris has a ~ 0.7 dex lower mean photometric metallicity and possibly higher intrinsic velocity dispersion than the GSS core. The GSS is also identified in the field of the M31 dwarf spheroidal satellite And I; the GSS in this field has a metallicity distribution identical to that of the GSS core. We confirm the previous finding of two kinematically cold components in stream C, and measure intrinsic velocity dispersions of ~ 10 and ~ 4 km s^{−1}. This compilation of the kinematical (mean velocity, intrinsic velocity dispersion) and chemical properties of stars in the GSS core and envelope, coupled with published surface-brightness measurements and wide-area star-count maps, will improve constraints on the orbit and internal structure of the dwarf satellite progenitor.

Key words: galaxies: halos – galaxies: individual (M31) – stars: kinematics – techniques: spectroscopic

Online-only material: color figures

1. INTRODUCTION

Galaxy mergers play a key role in galaxy formation and evolution (e.g., Searle & Zinn 1978; White & Rees 1978; Abraham et al. 1996; Springel et al. 2005). Observations have uncovered abundant evidence of the ubiquity of minor mergers in the form of tidal streams. Those discovered in our own Milky Way (MW) include the Magellanic stream (Mathewson et al. 1974), Monoceros stream (Yanny et al. 2003; Rocha-Pinto et al. 2003), and Sagittarius Stream (Ibata et al. 1994; Majewski et al. 2003; Newberg et al. 2003). Detailed studies of tidal debris features can provide insight into statistical models of hierarchical structure formation, the structure and merger history of an individual galaxy, and the properties of dwarf satellite systems. Due to their sparse stellar density and long dynamical times, stellar halos are ideal for investigating the merger history of an individual galaxy in detail as it is possible for higher density tidal debris features to remain identifiable in phase space for billions of years.

Observations of the Galaxy and theoretical arguments indicate that present-day dwarf satellite galaxies are different from the earlier generations of dwarf galaxies that merged to form the bulk of the virialized stellar halo. Stars in the classical satellites of the Galaxy on average have lower relative abundances of

alpha elements than stars of similar metallicities in the Galaxy’s halo (Fuhrmann 1998; Shetrone et al. 2001, 2003; Venn et al. 2004; Geisler et al. 2007), although the relative abundances of alpha elements are similar in the lowest metallicity stars (Kirby et al. 2009; Frebel et al. 2009). This has been shown to be a natural consequence of the hierarchical formation framework: simulations of stellar halo formation indicate that the majority of stellar mass in the halo was contributed by massive satellite galaxies that were accreted and disrupted early on, while surviving satellites are typically accreted much later (resulting in more metal enrichment via Type Ia supernovae) and are on average less massive (and experienced less efficient star formation) than the average halo building block (Robertson et al. 2005; Font et al. 2006a).

Minor mergers caught in the process of disruption bridge the gap between the surviving and completely disrupted dwarf populations. The giant southern stream (GSS; Ibata et al. 2001) in the Andromeda Galaxy (M31) offers a rare glimpse at a recent, significant accretion event in an external galaxy that is close enough ($d_{\text{M31}} = 783$ kpc; Stanek & Garnavich 1998; Holland 1998) to allow spectroscopy of individual red giant branch (RGB) stars. This stream is believed to result from the merger with M31 of a galaxy of stellar mass $\sim 2 \times 10^9 M_{\odot}$ less than 1 Gyr ago (Fardal et al. 2007), and is estimated to have

contributed at least 10% of the stars in M31’s halo (assuming azimuthal symmetry of M31’s halo; Font et al. 2006b; Ibata et al. 2007).

Ground-based photometric observations have shown that the GSS pollutes a large portion of the southern quadrant of M31’s stellar halo. It is observable to a projected distance of $R_{\text{proj}} \sim 100$ kpc from the center of M31, with a metal-rich “core”⁷ surrounded by a lower metallicity “envelope” (Ibata et al. 2007). The line-of-sight distance to the GSS has been measured from $R_{\text{proj}} \sim 10$ –60 kpc using the apparent magnitude of the tip of the RGB. The GSS is located ~ 100 kpc behind M31 at $R_{\text{proj}} \sim 60$ kpc, but is at the same distance as M31 at $R_{\text{proj}} \sim 10$ kpc (McConnachie et al. 2003). The detailed star formation history of a field located on the GSS has been derived from ultra-deep *Hubble Space Telescope* (*HST*) observations that reach below the main-sequence turnoff; the mean age of the stellar population is 8.8 Gyr (70% of the stars are younger than 10 Gyr), and the mean metallicity is -0.7 dex (Brown et al. 2006a).

Spectroscopic observations of RGB stars in M31’s GSS have been limited to the metal-rich core. Line-of-sight velocity distributions have been published for fields ranging from $R_{\text{proj}} = 12$ –60 kpc in projected radial distance from M31’s center (Ibata et al. 2004; Guhathakurta et al. 2006; Kalirai et al. 2006a). At $R_{\text{proj}} = 60$ kpc, the GSS’s mean line-of-sight (heliocentric) velocity is $v_{\text{hel}} = -320$ km s⁻¹. The innermost unambiguous detection of the GSS is in a field at $R_{\text{proj}} = 21$ kpc, with a mean line-of-sight velocity that is blueshifted substantially relative to the 60 kpc field: $v_{\text{hel}} = -513$ km s⁻¹ (Kalirai et al. 2006a). The 12 kpc field does not have a secure GSS detection; only 2–4 stars are found that may be related to the GSS based on their velocities (Ibata et al. 2004). In each field in which it has been observed, the GSS has a narrow velocity dispersion of ~ 10 –20 km s⁻¹.

These existing photometric and spectroscopic data sets have inspired detailed models of the collision that produced M31’s GSS. It is believed to have formed following a close pericentric passage of an M31 satellite on a highly radial orbit (Ibata et al. 2004; Font et al. 2006b; Geehan et al. 2006; Fardal et al. 2006, 2007; Mori & Rich 2008). The Fardal et al. (2007) numerical models showed that several other observed substructures in M31 (the northeastern (NE) and western (W) shelves) can be explained as the forward continuation of the GSS, and predict a third southeastern (SE) shelf feature that has since been discovered spectroscopically (Gilbert et al. 2007). The latest models demonstrate that several of the observed properties of the GSS and other debris in M31, including the asymmetric surface-brightness distribution of the GSS (McConnachie et al. 2003), the metal-rich core and metal-poor envelope structure (Ibata et al. 2007), and perhaps even the minor-axis arcs (Ibata et al. 2007), can be explained by a rotating progenitor with a reasonable initial metallicity gradient (Fardal et al. 2008). Despite these recent successes, the current location of the progenitor’s core and whether or not it is still intact remain unknown. More observations of the GSS and related debris are needed to further constrain models of the collision and produce more precise predictions for the current location and dynamical state of the progenitor’s core.

As part of the Spectroscopic and Photometric Landscape of Andromeda’s Stellar Halo (SPLASH) survey, we have con-

ducted an extensive spectroscopic survey of M31 RGB stars using the DEIMOS spectrograph on the Keck II telescope (Guhathakurta et al. 2005; Kalirai et al. 2006b; Gilbert et al. 2006, 2007), including observations of the GSS (Guhathakurta et al. 2006; Kalirai et al. 2006a). In this contribution, we present spectroscopic results for four new fields (and reanalysis of two older fields) in the vicinity of the GSS. We identify multiple kinematically cold features in their velocity distributions that we associate with the GSS or other photometric overdensities in M31’s stellar halo. These include the innermost unambiguous kinematical detection of the GSS and the first kinematical detection of GSS “envelope” debris. We also present the kinematical detection of substructure in the region of the photometrically identified 60 kpc minor-axis arc (Ibata et al. 2007), which some models of the collision suggest could be related to the GSS (Fardal et al. 2008). In addition, we further characterize the properties of a secondary kinematically cold component discovered by Kalirai et al. (2006a) in their study of a GSS core field at $R_{\text{proj}} = 21$ kpc. The kinematical and chemical properties of stars in the GSS presented here will provide improved observational constraints on the orbit and internal structure of the dwarf satellite progenitor and M31’s potential.

The observations, reduction procedures, and our technique for identifying a clean sample of M31 RGB stars are described in Section 2. The properties of stars in individual spectroscopic fields are discussed in Section 3: the kinematical components in each field are identified and their properties (mean velocity and velocity dispersion) are measured. The properties of the GSS and other debris features are discussed and compared in Section 4. Finally, the observations are discussed in the broader context of stellar halos in Section 5 and the main results of the paper are summarized in Section 6.

2. DATA

In this section, we describe the data reduction and analysis techniques that produce the final M31 RGB sample analyzed in the following sections. We describe the selection of fields from our spectroscopic survey (Section 2.1), give an overview of our photometric and spectroscopic observations and reductions (Sections 2.2 and 2.3), and discuss the [Fe/H] measurements used in our survey (Section 2.4). Finally, we describe the technique used to isolate the general M31 field population of RGB stars from MW dwarf stars and M31 dwarf spheroidal members (Section 2.5), and present the resulting velocity distributions of stars (Section 2.6).

2.1. Field Selection

The data presented here represent a small subset of the full data set collected as part of our SPLASH Keck/DEIMOS spectroscopic survey of M31 RGB stars. The subset was selected based on location of the fields with respect to M31’s GSS and other debris that may be related to the GSS (Figure 1). In relatively high-density regions of M31, “fields” consist of one spectroscopic mask (field f207) or two overlapping spectroscopic masks (field H13s), while in lower density regions fields consist of multiple masks designed from a single Kitt Peak National Observatory (KPNO) 4 m/Mosaic pointing (e.g., fields a13 and m4). Fields d1 and d3 were initially targeted to probe the stellar kinematics of the dwarf spheroidals (dSphs) And I and And III (J. S. Kalirai et al. 2009, in preparation). We are using them here to study the stellar kinematics of the underlying M31 field halo population, including possible GSS stars.

⁷ In this context “core” refers to the high-surface-brightness portion of the GSS extending along its length, not the GSS’s progenitor, which has yet to be identified.

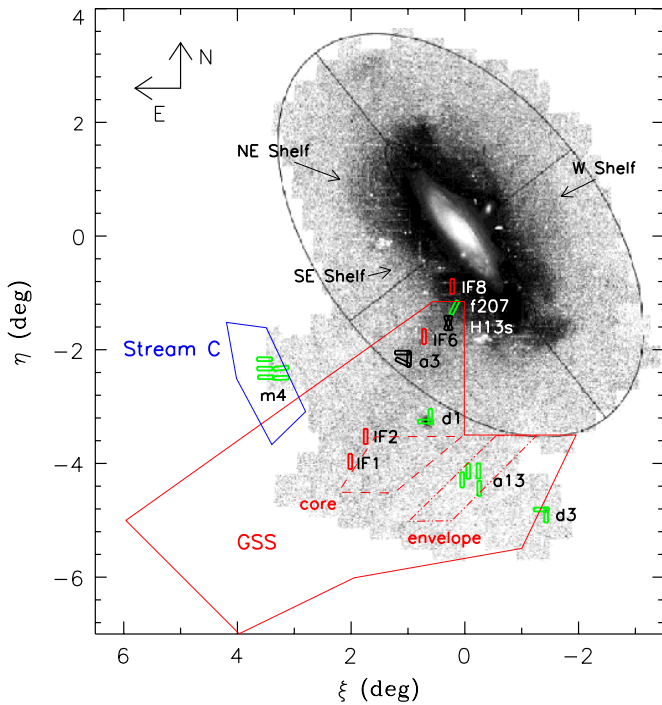


Figure 1. Locations (in M31-centric coordinates) of the Keck/DEIMOS spectroscopic fields discussed in this paper (Section 2), overlaid on a star-count map of M31 (Ibata et al. 2005). The GSS is the overdensity of stars extending to the south-southeast of M31's center. The approximate size and position angle of each DEIMOS spectroscopic slit mask are indicated by a rectangle. Black rectangles denote the location of masks from our survey of M31 for which velocity distributions have been previously published (a3 and H13s; Guhathakurta et al. 2006; Kalirai et al. 2006a), while green rectangles denote masks whose velocity distributions are being presented for the first time. Red rectangles denote masks IF1, IF2, IF6, and IF8 from Ibata et al. (2004). The polygons show the regions chosen by Ibata et al. (2007) for analysis of the photometric properties of the GSS (red solid outline), GSS core (red dashed outline), GSS envelope (red dash-dotted outline), and the 60 kpc southeast minor-axis arc “stream C” (blue solid outline).

Fields f207, H13s, and a3 are located on/near the highest-surface-brightness region of the GSS core (near its eastern edge) at $R_{\text{proj}} \sim 17$, 21, and 33 kpc, respectively. Fields d1 and a13 are located to the west of the highest-surface-brightness region of the GSS at $R_{\text{proj}} \sim 45$ and 58 kpc, respectively, but are still well within the broader region spanned by the GSS (Ibata et al. 2007), while field d3, at $R_{\text{proj}} \sim 69$ kpc, is on the western edge of this region. Finally, field m4 is located on M31's southeast minor axis at $R_{\text{proj}} \sim 57$ kpc, overlapping one of the minor-axis arc features (“stream C”) identified as stellar streams in the photometric survey of Ibata et al. (2007). The latest models show that for certain collision geometries, similar arc morphologies on the minor-axis can be produced, and thus stream C could be related to the GSS (Fardal et al. 2008).

2.2. Photometric and Spectroscopic Observations

Photometric catalogs for fields f207 and H13s were derived from images taken with the MegaCam instrument⁸ on the 3.6 m Canada–France–Hawaii Telescope (CFHT) (for details, see Kalirai et al. 2006a). Images were obtained with the g' and i' filters. The observed stellar magnitudes were transformed to Johnson–Cousins V and I magnitudes using observations

of Landolt photometric standard stars (Kalirai et al. 2006b). Photometry for the remainder of the fields (a3, a13, m4, d1, and d3) was derived from images taken with the Mosaic Camera on the KPNO 4 m telescope.⁹ The KPNO images were obtained with the Washington system M and T_2 filters and the intermediate-width DDO51 filter (Ostheimer 2003). The observed magnitudes were transformed to V and I magnitudes using the transformation equations of Majewski et al. (2000).

The DDO51 filter measures the surface-gravity sensitive Mg b and MgH stellar absorption features. Therefore, photometry in the intermediate-width DDO51 band, when combined with photometry in the M and T_2 bands, can be used as a discriminant of stellar surface gravity. This allows photometric preselection of spectroscopic targets that are likely to be RGB stars in M31 (as opposed to foreground MW dwarf stars). For masks designed from Mosaic photometry, the spectroscopic target selection algorithm gave highest priority to objects with high photometric RGB probability (as determined from the star's position in the $(M - \text{DDO51})$ versus $(M - T_2)$ color–color diagram (e.g., Palma et al. 2003; Majewski et al. 2005)), star-like morphology (based on the DAOPHOT parameters χ and sharp), and with I_0 magnitudes within ~ 1 mag of the tip of M31's RGB. The available space remaining on the mask was then packed with lower priority “filler” targets. All spectroscopic targets were restricted to the brightness range $20 < I_0 < 22.5$. In fields imaged with CFHT/MegaCam, spectroscopic targets were selected based on their I_0 magnitude and the SExtractor morphological criterion stellarity . Details of the spectroscopic target selection procedure for fields with and without DDO51 photometry are given by Guhathakurta et al. (2006) and Kalirai et al. (2006a), respectively.

Spectra were obtained using the DEIMOS spectrograph on the Keck II 10 m telescope, using the 1200 line mm^{-1} grating, which has a dispersion of $0.33 \text{ \AA pixel}^{-1}$, centered at $\sim 7800 \text{ \AA}$. The observations were obtained over five observing seasons (Fall 2002–2006) and are summarized in Table 1. The spectra were reduced using modified versions of the *spec2d* and *spec1d* software developed at the University of California, Berkeley in support of the DEEP2 galaxy redshift survey.¹⁰ These routines perform standard spectral reductions: flat-fielding, night-sky emission line removal, extraction of the two-dimensional spectra, and cross-correlation of the one-dimensional spectra with template stellar spectra. The template library includes stellar spectra obtained with DEIMOS and galaxy templates from the Sloan Digital Sky Survey. The cross-correlation procedure has been tested extensively against radial velocity standards (Simon & Geha 2007), and successfully reproduces known velocities. All spectra and cross-correlations were visually inspected, and only spectra with secure velocity measurements were included in further analysis. The observed velocity was corrected both for the motion of the Earth around the Sun (heliocentric correction) and for imperfect centering of the star within the slit (based on the observed position of the atmospheric A -band absorption feature relative to night-sky emission lines; Simon & Geha 2007; Sohn et al. 2007). The typical error in the velocity measurements is $\sim 6 \text{ km s}^{-1}$ for the M31 RGB stars presented here. It is calculated by adding in quadrature the

⁸ MegaPrime/MegaCam is a joint project of CFHT and CEA/DAPNIA, at the Canada–France–Hawaii Telescope which is operated by the National Research Council of Canada, the Institut National des Sciences de l'Univers of the Centre National de la Recherche Scientifique of France, and the University of Hawaii.

⁹ Kitt Peak National Observatory of the National Optical Astronomy Observatory is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

¹⁰ See <http://astro.berkeley.edu/~cooper/deep/spec2d/primer.html> and <http://astro.berkeley.edu/~cooper/deep/spec1d/primer.html>

Table 1
Details of Spectroscopic Observations and Basic Results

Mask	Projected Radius (kpc)	Pointing Center		P.A. (° E of N)	Date of Obs. (UT)	Photometric Bandpasses ^a	No. of Sci. Targets ^b	No. of Stars ^{b,c}	No. of M31 Stars ^{b,d}
		α_{J2000} (h:m:s)	δ_{J2000} (°:′:″)						
f207_1	17.4	00:43:42.64	+40:00:31.6	−27.0	2005 Nov 4	g', i'	202	123	114
H13s_1 ^e	21.3	00:44:14.76	+39:44:18.2	+21.0	2004 Sep 20	g', i'	134	100	87
H13s_2 ^e	21.3	00:44:14.76	+39:44:18.2	−21.0	2004 Sep 20	g', i'	138	121	98
a3_1 ^f	33.9	00:48:21.16	+39:02:39.2	+64.2	2002 Aug 16	$M, T_2, DDO51$	85	20	17
a3_2 ^f	32.5	00:47:47.24	+39:05:56.3	+178.2	2002 Oct 11	$M, T_2, DDO51$	81	32	28
a3_3 ^f	31.9	00:48:23.17	+39:12:38.5	−90.0	2003 Oct 26	$M, T_2, DDO51$	83	43	30
d1_1	44.1	00:45:48.61	+38:05:46.8	0.0	2005 Nov 5	$M, T_2, DDO51$	155	91	21
d1_2	45.6	00:46:13.95	+38:00:27.0	+90.0	2006 Sep 16	$M, T_2, DDO51$	149	96	28
a13_1	58.5	00:42:58.34	+36:59:19.3	0.0	2003 Sep 30	$M, T_2, DDO51$	80	25	10
a13_2	60.6	00:41:28.27	+36:50:19.2	0.0	2003 Sep 30	$M, T_2, DDO51$	71	23	10
a13_3	56.4	00:42:25.76	+37:08:25.4	0.0	2005 Nov 5	$M, T_2, DDO51$	78	29	13
a13_4	56.5	00:41:32.15	+37:08:37.4	0.0	2005 Nov 5	$M, T_2, DDO51$	86	32	11
d3_1	68.3	00:36:03.83	+36:27:27.4	+90.0	2005 Sep 8	$M, T_2, DDO51$	120	89	5
d3_2	69.8	00:35:39.61	+36:21:41.8	0.0	2005 Sep 8	$M, T_2, DDO51$	122	84	8
m4_1	59.3	01:00:40.24	+38:42:07.6	−90.0	2006 Sep 20	$M, T_2, DDO51$	72	30	13
m4_2	58.3	01:00:46.90	+38:51:05.9	−90.0	2006 Sep 21	$M, T_2, DDO51$	81	22	5
m4_3	57.1	01:00:50.72	+39:01:00.6	−90.0	2006 Nov 20	$M, T_2, DDO51$	90	40	15
m4_4	56.3	00:59:15.92	+38:42:12.0	−88.3	2006 Nov 20	$M, T_2, DDO51$	70	31	6
m4_5	54.9	00:59:18.67	+38:52:22.5	−84.0	2006 Nov 23	$M, T_2, DDO51$	75	38	14

Notes.

^a The catalogs used for spectroscopic target selection were derived from photometry obtained in the listed photometric bandpasses (Section 2.2).

^b A number of targets were observed on two different masks. The total number of *unique* science targets/stars/M31 RGB stars in fields H13s, a3, d1, and d3 is therefore less than the reported number. The number of stars with velocities recovered in two separate observations is 9 in H13s, 7 in a3, 3 in d1, and 11 in d3. Of these, 7 stars in H13s, 6 in a3, and 1 in d1 are M31 RGB field stars. Of the duplicates in the dwarf satellite fields, one each in d1 and d3 are MW dwarf stars; the remainder are dSph members.

^c The number of stellar spectra with secure velocity measurements.

^d The number of M31 RGB stars is defined as the number of stars that are identified as secure or marginal M31 RGB stars by the Gilbert et al. (2006) diagnostic method (Section 2.5). For the d1 and d3 masks, additional selection criteria have been applied to exclude members of the dSphs (Section 2.5.1).

^e Previously published in Kalirai et al. (2006a), the data presented here have undergone new reductions with our improved pipeline (Section 2.3). The sample of M31 RGB stars presented here is 70% larger than in Kalirai et al. (2006a).

^f Previously published in Guhathakurta et al. (2006), the data presented here have undergone new reductions with our improved pipeline (Section 2.3). The data were originally published without the benefit of the Gilbert et al. (2006) diagnostic method for determining M31 membership (Section 2.5).

random velocity measurement error (estimated from the cross-correlation routine) and a systematic error of 2.2 km s^{−1} (estimated via repeat observations of stars; Simon & Geha 2007).

2.3. Previously Published Fields

The fields H13s and a3 have been previously analyzed by Kalirai et al. (2006a) and Guhathakurta et al. (2006), respectively. These were among the first fields observed and reduced in our M31 spectroscopic survey. Since the original reduction of these fields, we have optimized our pipeline for reducing and measuring the velocities of stellar objects. It now includes a greater number of spectral templates (obtained with the same instrument as the science spectra) for measuring velocities via cross-correlation as well as a correction for imperfect centering of the star in the slit (Section 2.2). These improved data reduction and cross-correlation routines result in smaller systematic and random velocity errors. Velocities obtained with the improved cross-correlation procedure are shifted by $\sim +20$ km s^{−1} from the velocities obtained with the original cross-correlation pipeline. The typical measured velocity error of M31 RGB stars in these fields has been reduced from ~ 15 km s^{−1} to ~ 6 km s^{−1}.

The data in fields a3 and H13s presented here have been re-reduced using the new cross-correlation routine. All papers published by our group from 2007 onward utilize our improved

reduction pipeline. We have also recovered a significant fraction of spectra in field H13s that had failed in our original reduction pipeline. The sample of M31 RGB stars in field H13s is now 70% larger than the sample presented by Kalirai et al. (2006a).

In addition to the improvement in the reduction pipeline, the data in field a3 have benefited from an improved method to separate M31 RGB stars from foreground MW dwarf stars (Gilbert et al. 2006, see Section 2.5). This method had not yet been fully developed when these data were first presented. It is particularly important for securely identifying RGB stars in the kinematically broad spheroid of M31, whose radial velocities can overlap the tail of the MW dwarf star velocity distribution.

2.4. Photometric and Spectroscopic Metallicity Estimates

We estimate metallicities using both photometric and spectroscopic techniques for each star in our M31 survey. Like Kalirai et al. (2006b), we estimate photometric metallicities, $[\text{Fe}/\text{H}]_{\text{phot}}$, by comparing the star's position in the $[I_0, (V - I)_0]$ color-magnitude diagram (CMD) to a finely spaced grid of theoretical isochrones ($t = 12$ Gyr, $[\alpha/\text{Fe}] = 0.0$; VandenBerg et al. (2006)). The observed stellar magnitudes are transformed into Johnson–Cousins V and I magnitudes using the transformations described in Section 2.2, and extinction/reddening corrections based on the Schlegel et al. (1998) dust maps are applied on a

star-by-star basis. A distance modulus of $(m - M)_0 = 24.47$ (corresponding to a distance to M31 of 783 kpc; Stanek & Garnavich 1998) is applied to the isochrones. Although several systematic differences in our photometric metallicity estimates could exist between the data from the two different filter systems (Kalirai et al. 2006b), tests of the data show that any such differences are small compared to other sources of systematic error in our photometric metallicity estimates (discussed below).

The spectroscopic metallicity estimates, $[\text{Fe}/\text{H}]_{\text{spec}}$, are based on the strength of the Ca II-triplet absorption lines at $\lambda \sim 8500 \text{ \AA}$ using the empirical calibration relations of Rutledge et al. (1997a, 1997b) and assuming a horizontal branch magnitude of $I_0 = 25.17$ for M31 stars. The calibration is derived from RGB stars in MW globular clusters that span a range of metallicities.

It is reassuring that in general our photometric metallicity estimates are in good agreement with the spectroscopic metallicity estimates based on the strength of the Ca II absorption line triplet. The individual Ca II equivalent width (EW) measurements and derived $[\text{Fe}/\text{H}]_{\text{spec}}$ estimates have large random errors (Gilbert et al. 2006) due to the low typical signal to noise of our M31 RGB spectra ($S/N \sim 15 \text{ \AA}^{-1}$). Since the $[\text{Fe}/\text{H}]_{\text{spec}}$ estimates are very noisy, we can only make a meaningful comparison if we bin our sample of stars in terms of $[\text{Fe}/\text{H}]_{\text{phot}}$ and compute the average $[\text{Fe}/\text{H}]_{\text{spec}}$ in each bin (Figure 7 of Kalirai et al. 2006b). We thus use photometric rather than spectroscopic metallicity estimates for the analysis presented here.

Although the random error is smaller for $[\text{Fe}/\text{H}]_{\text{phot}}$ than $[\text{Fe}/\text{H}]_{\text{spec}}$, there are substantial systematic errors in $[\text{Fe}/\text{H}]_{\text{phot}}$ due to differences between the age and alpha enrichment of a star and the adopted isochrones. Moreover, there are substantial variations between the isochrones computed by different groups. Photometric metallicity estimates based on the Padova (Girardi et al. 2002) and Yale–Yonsei (Y^2 ; Demarque et al. 2004) isochrone sets vary from those based on the Victoria–Regina (VR; VandenBerg et al. 2006) isochrones by a median amount of ~ 0.1 dex. The variation across the isochrone sets is largest for the bluest stars (~ 0.3 – 0.5 dex).

There are no observational constraints on the alpha enrichment of M31 halo stars, but the assumption of a single age for the stellar population is known to be incorrect. Deep *HST*/Advanced Camera for Surveys (ACS) observations of fields in M31's stellar halo and the GSS that reach below the main-sequence turnoff have shown that the age of the stellar population ranges from ~ 8 to 13 Gyr in the halo and ~ 6 to 13 Gyr in the GSS; $\sim 33\%$ of the population is younger than 10 Gyr in M31's halo, while $\sim 70\%$ of the population is younger than 10 Gyr in the GSS field (Brown et al. 2006a, 2007). These studies find that in general the intermediate-age stars are relatively metal rich ($-0.75 \lesssim [\text{Fe}/\text{H}] \lesssim 0.4$) while the older populations are relatively metal poor ($[\text{Fe}/\text{H}] \lesssim -1.0$).

Our photometric metallicity estimates are therefore likely to be biased low for the most metal-rich stars by our assumption of an old (12 Gyr), coeval stellar population, since the *HST* studies show that these metal-rich stars are typically of intermediate age. Changing the isochrone age from 12 to 6 Gyr results in a $\sim +0.15$ dex shift in the derived stellar metallicity for stars with $[\text{Fe}/\text{H}]_{\text{phot}} > -1.0$, and a $\sim +0.3$ dex shift for stars with $-2.3 < [\text{Fe}/\text{H}]_{\text{phot}} < -1.5$. Although the systematic bias in the derived metallicity estimates introduced by assuming an old stellar population is large for intermediate-age, metal-poor stars, the *HST* studies show that such stars are rare in both M31's halo and the GSS (Brown et al. 2006a).

2.5. Isolation of the Field M31 RGB Sample

Even in fields with DDO51-based photometric preselection of likely M31 RGB stars (Section 2.2), MW dwarf stars along the line of sight to M31 still constitute a significant fraction of the recovered stellar spectra. Since the radial velocity distributions of MW dwarf stars and M31 RGB stars overlap, we use a set of photometric and spectroscopic diagnostics to determine the probability that an individual star is a red giant at the distance of M31 or a dwarf star in the MW. The diagnostics are (1) radial velocity, (2) photometry in the Washington M and T_2 bands and the DDO51 band, (3) Na I EW versus $(V - I)_0$ color, (4) position in an $[I_0, (V - I)_0]$ CMD, and (5) spectroscopic versus photometric $[\text{Fe}/\text{H}]$ estimates. This technique for separating the M31 RGB and MW dwarf star populations is presented and tested in detail by Gilbert et al. (2006). We briefly summarize it here.

Each of the five diagnostics provides separation between M31 RGB stars and MW dwarf stars based on different physical parameters. The combination of the five diagnostics provides a powerful test of whether an individual star is an M31 red giant or MW dwarf. While M31 RGB stars have a broad velocity distribution around M31's systemic velocity of $v_{\text{sys}} = -300 \text{ km s}^{-1}$, MW dwarf stars tend to share the Sun's rotation about the MW's center and have $v_{\text{helio}} \approx 0 \text{ km s}^{-1}$. As discussed in Section 2.2, the combination of the M , T_2 , and DDO51 photometry is sensitive to the difference in surface gravity between RGB and dwarf stars. The strength of the Na I doublet at $\lambda \sim 8190 \text{ \AA}$ is both surface-gravity and temperature sensitive. It provides good separation of RGB and dwarf stars with $(V - I)_0 \gtrsim 2$. M31 RGB stars follow the RGB isochrones in the CMD, while MW dwarf stars are more broadly distributed in CMD space. Finally, due to the fact that both the $[\text{Fe}/\text{H}]_{\text{phot}}$ and $[\text{Fe}/\text{H}]_{\text{spec}}$ estimates are based on the assumption that the star is an M31 RGB star, these estimates agree only if the star is in fact an RGB star at the distance of M31.

A star i is given a likelihood in a given diagnostic j of being an M31 RGB star, defined as $L_{ij} = \log(P_{\text{giant}}/P_{\text{dwarf}})$, where P is the probability the star is an RGB or dwarf star based on a comparison of the star's position in that diagnostic space to probability distribution functions based on empirical training sets of definite M31 RGB and MW dwarf stars (Gilbert et al. 2006, Section 3). The individual likelihoods are combined to form an overall likelihood $\langle L_i \rangle$. If $\langle L_i \rangle > 0$, it is more probable that the star is an RGB star than a dwarf star, and vice versa for $\langle L_i \rangle < 0$. Figure 2 shows the overall likelihood distributions of stars in each of our fields. For the purposes of the current analysis, we use stars with $\langle L_i \rangle > 0$ —i.e., more likely to be an M31 RGB star than an MW dwarf star. Furthermore, we include only stars whose positions in the $[I_0, (V - I)_0]$ CMD are within the isochrone grid ($[\text{Fe}/\text{H}]_{\text{phot}} = -2.31$ to 0.0) to within the average photometric error of the sample.¹¹

It is possible that stars with $\langle L_i \rangle > 0$ that are significantly bluer than the isochrone grid in the CMD may be asymptotic giant branch stars at the distance of M31 or M31 RGB stars with flawed photometry. However, MW dwarf stars along the line of sight with large negative velocities ($v_{\text{los}} \lesssim -150$) tend to have blue $(V - I)_0$ colors, as these stars are generally main-sequence turnoff stars in the MW's halo. Since several of the diagnostics have limited differentiating power between RGB and dwarf stars

¹¹ In the terminology of Gilbert et al. (2006), we include class +1, +2, and +3 stars in the following analysis, and exclude class -2 stars from our M31 RGB sample.

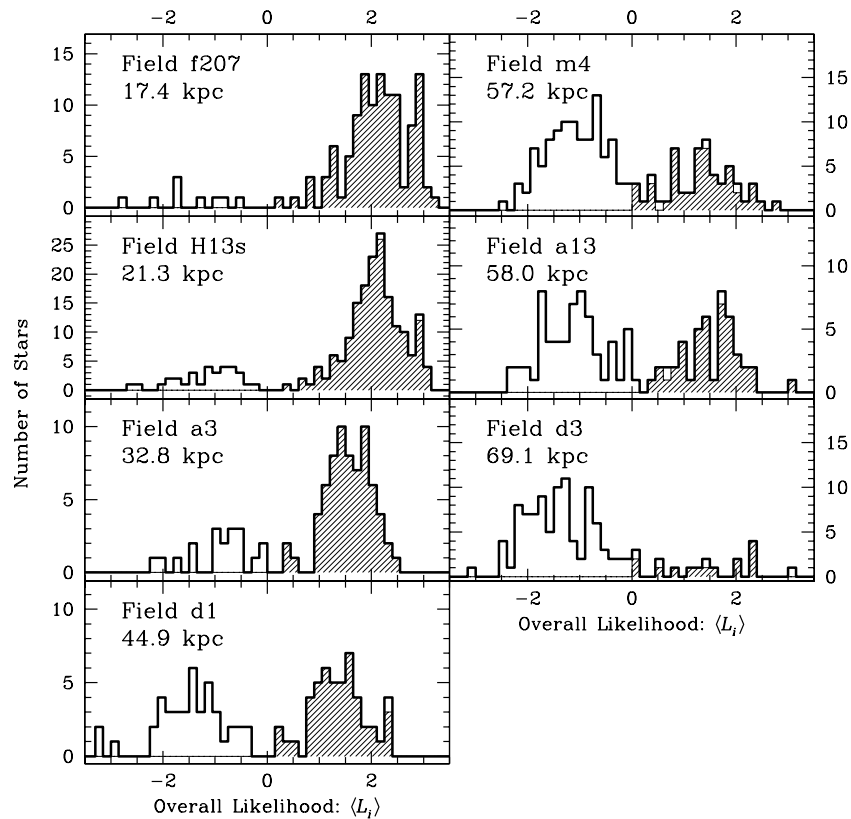


Figure 2. Overall likelihood distributions for stars in each of the fields analyzed in this paper (Section 2.5). The shaded histograms denote stars identified as M31 RGB stars. Stars with $\langle L_i \rangle > 0$ that are not shaded are stars that, despite having physical characteristics that are consistent with the M31 RGB population, have been removed from the M31 RGB sample due to their very blue $(V - I)_0$ colors. In the d1 and d3 fields, only stars *not* associated with the dSph galaxies are included in this distribution (Section 2.5.1).

for very blue stars [$(V - I)_0 \lesssim -1.2$], the $\langle L_i \rangle$ values of these stars are strongly influenced by their M31 RGB-like velocities (Gilbert et al. 2006, Section 4.1.2). We therefore choose to be conservative and exclude these very blue stars from our M31 RGB sample. These stars represent only 2% of the total stellar sample.

In Gilbert et al. (2007), we used a large sample of RGB stars in M31’s inner spheroid ($R_{\text{proj}} \lesssim 30$ kpc) to investigate the effect of removing the velocity diagnostic from our selection of M31 RGB and MW dwarf stars. This sample was ideal for this analysis since the density of RGB stars in M31’s inner regions is much higher than the density of MW stars along the line of sight. We found that while the exclusion of the velocity diagnostic does increase the completeness of the M31 RGB sample at velocities between -100 and 0 km s $^{-1}$, it also increases the level of MW dwarf star contamination at these velocities. Some of the fields presented here are at large distances from M31’s center ($R_{\text{proj}} \sim 60$ kpc), where the density of MW dwarf stars along the line of sight is significantly higher than the density of M31 RGB stars. We thus choose lower contamination over higher completeness for the present work, and include the velocity diagnostic in our selection of M31 RGB stars. This affects only stars with velocities in the range of the MW dwarf distribution. Since we do not fit the velocity distribution of M31’s spheroid in this work and the kinematically cold substructures presented below have mean velocities $\lesssim -250$ km s $^{-1}$ (Section 3), incompleteness of the M31 RGB sample at velocities near zero will only affect the estimated fraction of the stars in M31’s kinematically broad spheroid. Furthermore, this effect will be small since the portion of velocity space where we expect incompleteness due to the use

of the velocity diagnostic is both in the tail of the M31 spheroid velocity distribution and is small compared to the full velocity space used in the fits to the velocity distributions in Section 3.

2.5.1. Andromeda Satellite Fields

The dSph And I is located within the area spanned by the GSS, while the dSph And III is located just outside this region (Figure 1, marked by the “d1” and “d3” spectroscopic masks, respectively). In addition to dSph RGB stars and MW dwarf stars, M31 “field” stars (RGB stars that are physically associated with the halo of M31 rather than with the dSph) are expected to be present in our spectroscopic sample. The half-light radii of the dSphs And I and And III are $2'8$ (600 pc) and $1'6$ (360 pc), respectively (McConnachie & Irwin 2006), therefore much of the area of each $4' \times 16'$ DEIMOS slit mask primarily targets the M31 field star population.

M31 field RGB stars are expected to have different velocity, metallicity, and spatial distributions than the dSph members. Figures 3(a) and 4(a) display the velocity distributions of all stars observed in the d1 and d3 fields, respectively (thin open histograms). The bold shaded histograms denote the stars identified as red giants using the method described in Section 2.5. Although the majority of the RGB stars in the two dSph fields are part of the prominent kinematically cold velocity distributions of the And I and And III dSphs (at $v \sim -375$ and -345 km s $^{-1}$, respectively), there are also RGB stars spanning a wide range of velocities in each field, as expected for stars that are associated with M31’s spheroid.

We utilize the spatial, kinematic, and chemical properties of the RGB stars to identify individual RGB stars as likely M31 field stars. The spatial distributions of stars observed in the d1

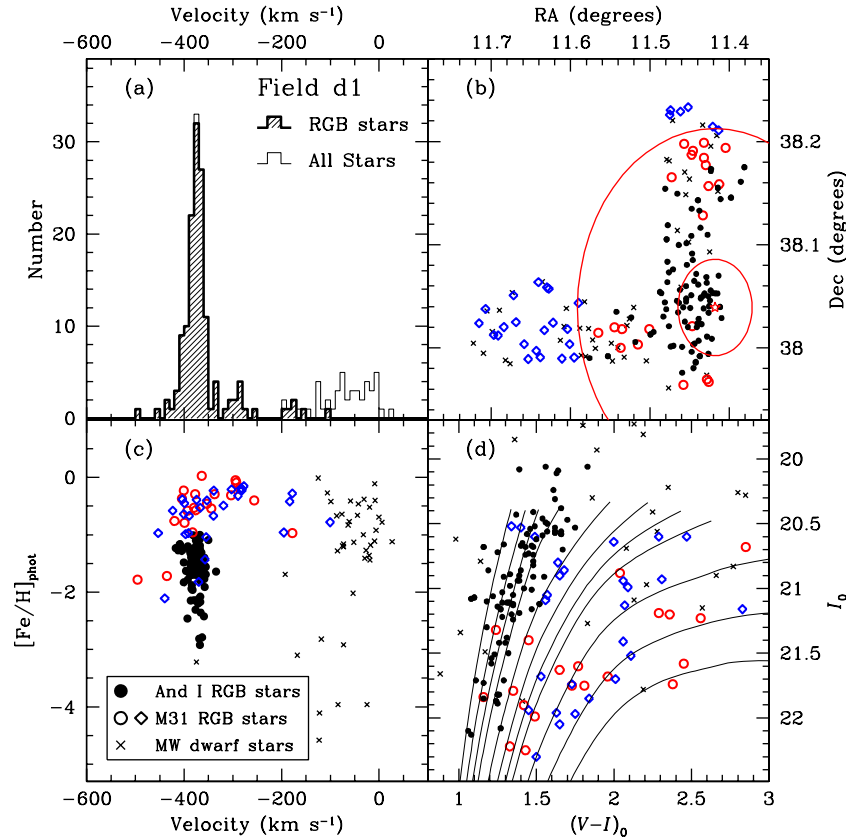


Figure 3. M31 field star selection in the d1 spectroscopic field (Section 2.5.1). (a) Line-of-sight velocity distribution of all stars (thin open histogram), and stars identified as RGB stars by the likelihood method described in Section 2.5 (thick, shaded histogram). In addition to the prominent kinematically cold peak formed by dSph members, there are RGB stars present in this field with a wide range of velocities, as expected for stars associated with M31's spheroid. (b) Sky positions of the RGB stars in field d1. The center (red star), half-light radius (inner ellipse), and tidal radius (outer ellipse) of the And I dSph are marked (McConnachie & Irwin 2006). RGB stars at or beyond the tidal radius of the dSph are designated as M31 field stars (open diamonds). Open circles denote red giants designated as M31 field stars by virtue of their separation from the locus of dSph members in the v - $[\text{Fe}/\text{H}]_{\text{phot}}$ plane (panel (c)). Filled circles denote RGB stars that do not meet our selection criteria for M31 field stars. The vast majority of these stars are dSph members. Crosses denote MW dwarf stars. (c) Velocity vs. $[\text{Fe}/\text{H}]_{\text{phot}}$ for all stars. RGB stars associated with the dSph form a tight locus in this space. A second population of RGB stars with higher $[\text{Fe}/\text{H}]_{\text{phot}}$ values ($[\text{Fe}/\text{H}]_{\text{phot}} > -1.0$) and a wider spread in velocities, most of which are beyond the tidal radius of the dSph, are more likely associated with M31's spheroid. (d) Location of M31 field stars and dSph members in the $[I_0, (V-I_0)]$ CMD. Theoretical isochrones shifted to the distance of M31 have been overlaid on the data. The isochrones are for an age of 12 Gyr, $[\alpha/\text{Fe}] = 0$, and span a range of $[\text{Fe}/\text{H}]$ from -2.31 to 0.0 (VandenBerg et al. 2006).

(A color version of this figure is available in the online journal.)

and d3 fields are displayed in Figures 3(b) and 4(b). The center of each dSph is marked by an open star. The inner and outer ellipses denote the half-light radii and tidal radii, respectively, of each dSph (McConnachie & Irwin 2006). All RGB stars beyond the tidal radius of the dSph are designated as M31 field stars (open diamonds). It is possible that some of these stars were tidally stripped from the dSph, in which case they are recent additions to M31's spheroid.

Figures 3(c)-(d) and 4(c)-(d) demonstrate that stars associated with the dSph galaxies (filled circles) form fairly tight loci in the v - $[\text{Fe}/\text{H}]_{\text{phot}}$ plane and the $[I_0, (V-I_0)]$ CMD. A small fraction of the RGB stars within the tidal radii of the dSphs are designated as likely M31 RGB stars by virtue of having velocities and/or $[\text{Fe}/\text{H}]_{\text{phot}}$ estimates that are well removed from the dSph locus (open circles). All but one of these stars are beyond the half-light radius of the dSph (Figures 3(b) and 4(b)). In the d3 field, these stars are removed from the mean velocity of the dSph by more than eight times the measured width of the dSph velocity distribution (σ_v). The RGB star in the d1 field at $v = -438 \text{ km s}^{-1}$ that is within the tidal radius of the dSph is $5.8\sigma_v$ removed from the mean velocity of And I. The full RGB population in the d1 field appears to be comprised primarily of two distinct populations: a relatively

metal-poor ($[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$) population spanning a narrow velocity range (the And I dSph), and a relatively metal-rich ($[\text{Fe}/\text{H}]_{\text{phot}} > -1.0$) population with a broad range of velocities. The stars beyond the tidal radius of And I comprise the majority of the more metal-rich population (Figures 3(c) and (d)), indicating that this population is not associated with the dSph. We therefore designate all RGB stars in d1 with $[\text{Fe}/\text{H}]_{\text{phot}} > -1.0$ as likely M31 field stars.

The selection criteria for M31 field stars described above yield an incomplete sample, as M31 field stars that fall both within the tidal radius of the dSph and near the region of v - $[\text{Fe}/\text{H}]_{\text{phot}}$ space occupied by the dSph members are not included in our sample. However, the fraction of M31 field stars that satisfy both these criteria is expected to be small, since (1) dSph members span a very narrow range in velocity and (2) the surface brightness of each dSph becomes increasingly dominant over M31's spheroid as one moves toward the center of the dSphs, resulting in an increasingly greater likelihood that spectroscopic slits are placed on a dSph member rather than an M31 field star. Only the M31 field stars will be analyzed in the current work. The properties of the dSph members in these fields will be discussed in a future paper (J. S. Kalirai et al. 2009, in preparation).

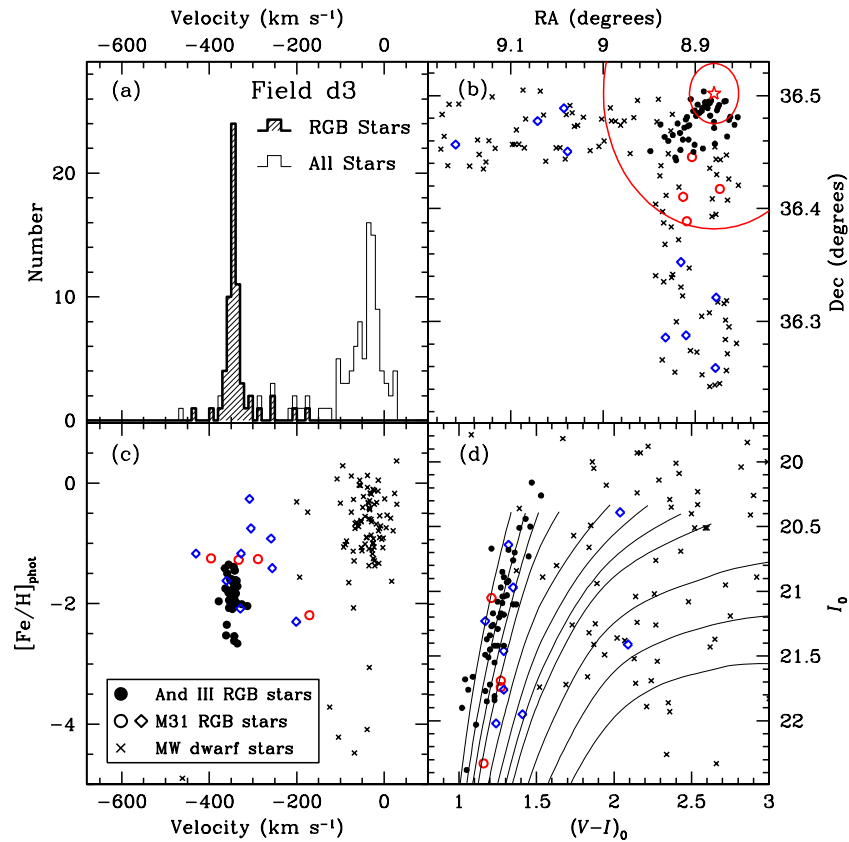


Figure 4. Same as Figure 3, for field d3 located on the dSph And III.
(A color version of this figure is available in the online journal.)

2.6. Velocity Distributions

Figure 5 presents the line-of-sight heliocentric velocity distributions of all stars (M31 red giants and MW dwarfs) observed in our fields, as well as the velocity distributions of stars in the four GSS spectroscopic fields published by Ibata et al. (2004; fields IF1, IF2, IF6, and IF8). In the innermost fields, M31 RGB stars greatly outnumber MW dwarf stars along the line of sight, while the majority of stars observed in the outermost fields are MW dwarf stars due to the decreasing density of M31's halo with R_{proj} .

Most of the fields show clear evidence of kinematically cold features in their stellar velocity distributions in addition to stars with a wide range of velocities (M31's kinematically broad spheroid). In fact, several fields appear to have multiple kinematically cold components (e.g., f207, H13s, and m4). The GSS is a prominent structure in the stellar velocity distributions of fields near the eastern edge of the high-surface-brightness GSS core (f207, H13s, IF6, a3, IF2, and IF1; Figure 1). The detection of the GSS in field f207 is the innermost secure kinematical detection of the GSS to date. There are also prominent kinematically cold peaks in fields d1 and a13 with velocities in keeping with the expected velocity of the GSS. We will show below that these are also attributable to the GSS, and are the first kinematical detections of the GSS in fields to the west of its prominent, eastern edge (Sections 3.3.1 and 4.1).

3. IDENTIFICATION OF SUBSTRUCTURE IN INDIVIDUAL FIELDS

This section discusses in detail the stellar velocity distributions and evidence for the presence of kinematically cold

components in each of our Keck/DEIMOS spectroscopic fields (Figure 5). The velocity and metallicity distributions of M31 RGB stars are analyzed in order to identify substructure in each field, and maximum-likelihood fits to the velocity distributions are performed to characterize the kinematical properties of the substructure, including its mean velocity, velocity dispersion, and the fraction of stars belonging to it. The fields near the eastern edge of the GSS are discussed first (f207, H13s, and a3; Section 3.2), followed by the fields to the west (d1, a13, and d3; Section 3.3.1) and the field on the 60 kpc minor-axis arc (m4; Section 3.3.2). A broader analysis of the kinematical and chemical properties of the identified substructures in relation to each other and a discussion of the origins of the features will be presented in Section 4.

3.1. Characterizing Kinematical Substructure

We use a combination of kinematics and metallicities to identify the number of kinematically cold substructures in each field. We then fit a multiple-component Gaussian to the kinematical distribution using a maximum-likelihood technique. The maximum-likelihood multi-Gaussian fits in each field are discussed below and summarized in Table 2. Quoted errors are 90% confidence limits from the maximum-likelihood analysis. Although the true shapes of the velocity distributions of the various kinematical components in M31's spheroid are likely different from pure Gaussians, they provide a convenient means of characterizing the mean velocity and velocity dispersions of kinematical features.

We assume a kinematically broad distribution of velocities corresponding to M31's well-mixed spheroid is present in each

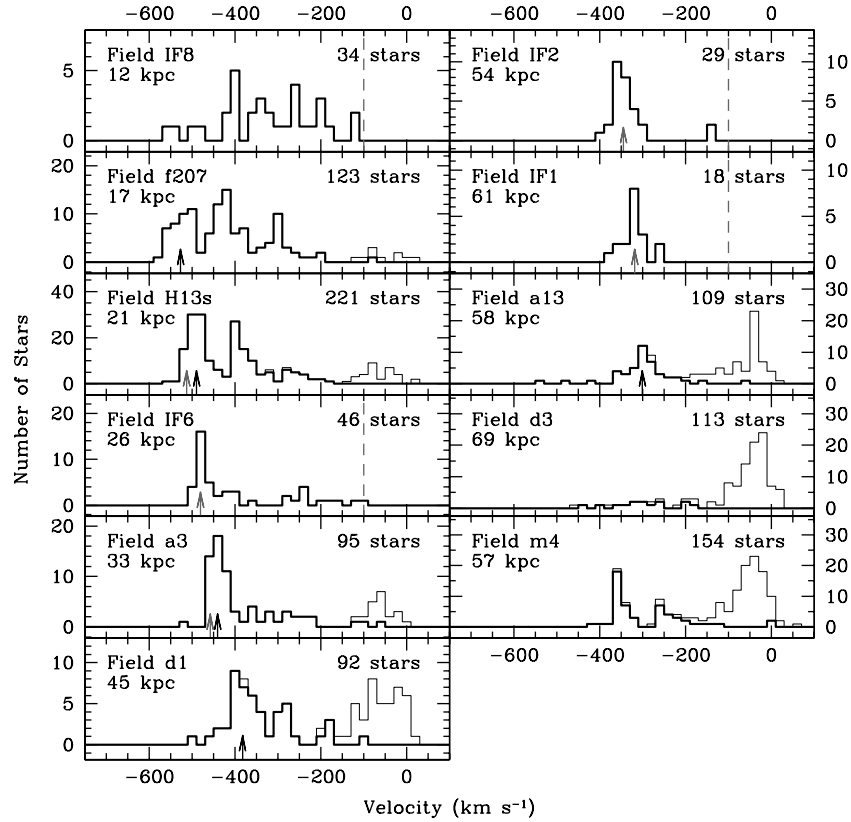


Figure 5. Line-of-sight, heliocentric velocity distributions for stars in each of the fields shown in Figure 1, displayed in order of increasing distance along the giant southern stream (GSS). The velocity distributions of stars identified as M31 RGB stars (Section 2.5) are shown by the bold histograms. The velocity distributions of all observed stars (including stars identified as MW dwarfs) are shown as thin histograms. Stars associated with the dSphs And I and And III have been removed from the velocity distributions. Arrows mark the mean velocity of the GSS: black arrows denote velocities measured in this work (Section 3), while gray arrows denote previously published velocities. The mean velocity of the GSS becomes increasingly negative as the stream approaches the center of M31. In addition to the prominent kinematical peaks identified as the GSS, several of our spectroscopic fields show evidence of other kinematical substructures (fields f207, H13s, and m4; Section 3). Each panel is labeled with the projected radial distance of the field from M31's center and the total number of stars with radial velocity measurements. The majority of the data come from the SPLASH Keck/DEIMOS survey of M31's stellar halo (fields f207, H13s, a3, d1, a13, m4, and d3). The velocities in fields IF8, IF6, IF2, and IF1 were obtained from Ibata et al. (2004); only stars with $v < -100$ km s⁻¹ are available from the literature (dashed lines). M31's systemic velocity is $v_{\text{sys}} = -300$ km s⁻¹.

Table 2
Results of Maximum-Likelihood Fits to the Velocity Distributions: Fields with One or More Kinematically Cold Component

Field	Primary Cold Component ^a			Secondary Cold Component ^b			Field Location
	$\langle v \rangle_{\text{primary}}$ (km s ⁻¹)	$\sigma_v^{\text{primary}}$ (km s ⁻¹)	$N_{\text{primary}}/N_{\text{total}}$	$\langle v \rangle_{\text{secondary}}$ (km s ⁻¹)	$\sigma_v^{\text{secondary}}$ (km s ⁻¹)	$N_{\text{secondary}}/N_{\text{total}}$	
f207	$-524.4^{+7.7}_{-7.4}$	$23.2^{+7.2}_{-5.0}$	$0.31^{+0.08}_{-0.07}$	$-425.6^{+10.0}_{-8.3}$	$20.8^{+12.9}_{-7.6}$	$0.31^{+0.11}_{-0.10}$	GSS core
H13s	$-489.7^{+4.5}_{-4.2}$	$21.3^{+4.0}_{-3.2}$	$0.48^{+0.07}_{-0.06}$	$-388.3^{+5.4}_{-5.4}$	$17.0^{+6.6}_{-6.9}$	$0.27^{+0.07}_{-0.08}$	GSS core
a3	$-440.5^{+5.2}_{-4.9}$	$16.8^{+4.6}_{-3.3}$	$0.59^{+0.11}_{-0.12}$	GSS core
d1	$-382.0^{+12.4}_{-11.6}$	$30.3^{+13.3}_{-8.6}$	$0.57^{+0.15}_{-0.17}$	$-289.2^{+6.8}_{-7.4}$	$8.2^{+8.6}_{-4.5}$	$0.15^{+0.11}_{-0.09}$	West of highest-surface-brightness region of GSS core
a13	$-300.7^{+12.0}_{-11.6}$	$32.2^{+11.8}_{-10.6}$	$0.72^{+0.15}_{-0.21}$	GSS envelope
m4	$-355.3^{+5.1}_{-4.5}$	$11.4^{+5.2}_{-4.1}$	$0.45^{+0.14}_{-0.13}$	$-254.8^{+5.9}_{-4.9}$	$6.6^{+6.0}_{-2.9}$	$0.16^{+0.10}_{-0.08}$	60 kpc minor-axis arc (stream C)

Notes.

^a Quoted errors are the 90% confidence limits from the maximum-likelihood analysis. All fields are assumed to also have a broad underlying distribution of stars with parameters $\langle v \rangle^{\text{sph}} = -300$ km s⁻¹ and $\sigma_v^{\text{sph}} = 129$ km s⁻¹ (Gilbert et al. 2007).

^b In addition to the primary kinematically cold component, fields f207, H13s, d1, and m4 contain second kinematically cold components.

field. To date, only two spectroscopic studies have measured the velocity dispersion of M31's spheroid. Gilbert et al. (2007) used a sample of ~ 1000 M31 RGB stars in spectroscopic fields from $R_{\text{proj}} = 9$ to 30 kpc along M31's minor axis. They measured a velocity dispersion of $\sigma_v^{\text{sph}} = 129$ km s⁻¹ for M31's spheroid,

and saw no evidence for a change in the velocity dispersion with radius between 9 and 30 kpc. The stars were selected using the technique described in Section 2.5, but without use of the velocity diagnostic. Chapman et al. (2006) measured an average velocity dispersion of $\sigma_v^{\text{sph}} = 126$ km s⁻¹ for M31's

spheroid, consistent with the Gilbert et al. (2007) value. Their measurement was based on a sample of ~ 800 M31 stars selected using windows in line-of-sight velocity space designed to remove foreground MW stars, M31 disk stars, and substructure. They also fit a model in which the velocity dispersion of M31's spheroid decreases with increasing projected distance. Using this model, they find that M31's spheroid is expected to have a velocity dispersion of $\sigma_v^{\text{sph}} = 137 \text{ km s}^{-1}$ at $R_{\text{proj}} = 17 \text{ kpc}$, decreasing to $\sigma_v^{\text{sph}} = 100 \text{ km s}^{-1}$ at $R_{\text{proj}} = 60 \text{ kpc}$.

Since the number of M31 spheroid stars at large radii ($\geq 40 \text{ kpc}$) in the (Chapman et al. 2006) study is small and a significant portion of velocity space was not used for measuring the change in the velocity dispersion with radius, we assume a constant value for the dispersion of M31's spheroid in analyzing the velocity distribution of our fields ($\langle v \rangle^{\text{sph}} = -300 \text{ km s}^{-1}$, $\sigma_v^{\text{sph}} = 129 \text{ km s}^{-1}$). However, for the fields at larger radii ($R_{\text{proj}} > 40 \text{ kpc}$), we will comment on the effect on the fit of assuming a smaller velocity dispersion for M31's spheroid.

3.2. GSS Core Fields

3.2.1. New GSS Field at $R \sim 17 \text{ kpc}$

Field f207 is located on the high-surface-brightness core of the GSS at a projected radius of 17 kpc from the center of M31 (Figure 1). The line-of-sight velocity distribution of stars in this field is shown in Figure 6(a), while panels (b)–(d) show the velocity distribution of stars in field f207 as a function of $[\text{Fe}/\text{H}]_{\text{phot}}$ (Section 2.4). The GSS is the kinematically cold, relatively metal-rich group of stars at very negative velocities ($v < -480 \text{ km s}^{-1}$). This is consistent with the expected velocity of the GSS at this position based on previous kinematical studies (Figure 5; Ibata et al. 2004; Kalirai et al. 2006a; Guhathakurta et al. 2006). The data in this field provide the innermost secure kinematical detection of the GSS to date.

In addition to the GSS, a second prominent kinematically cold, metal-rich population can be seen in the velocity range $-460 \text{ km s}^{-1} \lesssim v \lesssim -400 \text{ km s}^{-1}$. A less prominent, kinematically cold, primarily metal-rich group of stars may be present at $v \sim -300 \text{ km s}^{-1}$. As expected for a field at this projected radial distance from M31, the majority of the stars not in the kinematically cold peaks are also relatively metal rich (e.g., Durrell et al. 2001, 2004; Kalirai et al. 2006b; Gilbert et al. 2007).

The data are consistent (via a reduced- χ^2 test) with being drawn from the maximum-likelihood Gaussian with three kinematical components: the GSS core at $v = -524 \text{ km s}^{-1}$, a second kinematically cold component at $v = -426 \text{ km s}^{-1}$, and M31's kinematically broad spheroid (Table 2; Figure 6(a), solid curve). We therefore adopt the triple-Gaussian fit for this field in the analysis presented here. Future observations may increase the statistical significance of the $v \sim -300 \text{ km s}^{-1}$ component, which currently has an excess of only ~ 6 stars ($\sim 5\%$ of the M31 RGB stars in this field) above the number expected from M31's broad spheroid. The stars in this possible kinematical peak have a mean velocity of $v \sim -305 \text{ km s}^{-1}$ and a dispersion of $\sim 8 \text{ km s}^{-1}$.

The measured velocity dispersion of the GSS core and the second kinematically cold component are $\sigma_v^{\text{GSS}} = 23.2_{-5.0}^{+7.2} \text{ km s}^{-1}$ and $\sigma_v^{\text{sec}} = 20.8_{-7.6}^{+12.9} \text{ km s}^{-1}$, respectively. The median velocity error of M31 RGB stars in this field is 6.9 km s^{-1} (Section 2.2). Thus, the intrinsic velocity dispersions of the GSS and second kinematically cold component are $\sigma_v^{\text{intr}} = 22.2$ and $\sigma_v^{\text{intr}} = 19.6 \text{ km s}^{-1}$, respectively. We will present evidence

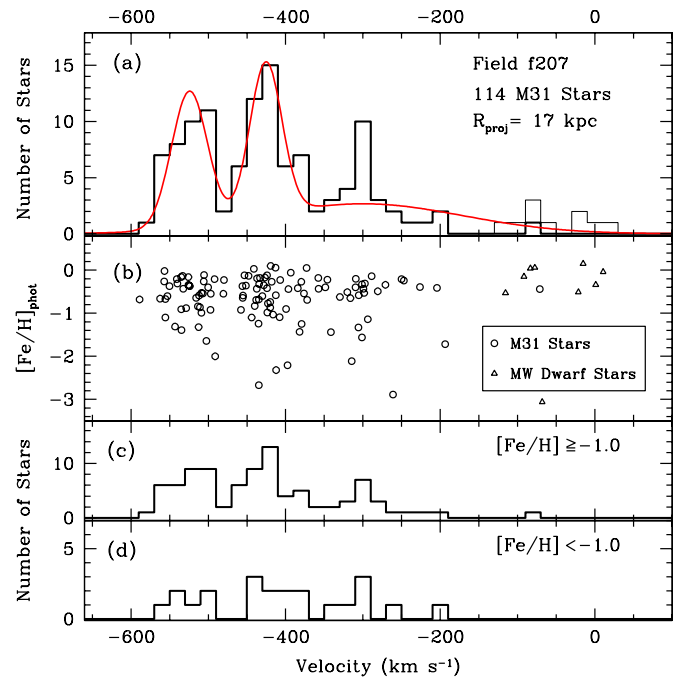


Figure 6. Line-of-sight, heliocentric velocities of stars in the GSS core field at $R_{\text{proj}} = 17 \text{ kpc}$ (field f207; Section 3.2.1). (a) The bold histogram denotes the distribution of stars classified as M31 red giants (Section 2.5). The red curve is the maximum-likelihood multi-Gaussian fit to the M31 RGB stars (Table 2). The thin histogram denotes the velocity distribution of all stars (M31 RGB and MW dwarfs). (b) Velocity vs. metallicity ($[\text{Fe}/\text{H}]_{\text{phot}}$) of the stars in field f207. Stars classified as M31 RGB stars are denoted by circles and stars classified as MW dwarf stars are denoted by triangles. (c) Velocity distribution of M31 RGB stars with $[\text{Fe}/\text{H}]_{\text{phot}} \geq -1.0$. (d) Velocity distribution of M31 RGB stars with $[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$. The GSS core is evident as a kinematically cold peak of relatively metal-rich stars at $v \sim -525 \text{ km s}^{-1}$ in panels (a)–(c), and a second kinematically cold, relatively metal-rich group of stars is seen at $v \sim -425 \text{ km s}^{-1}$. In addition, stars associated with M31's spheroid are present with a wide range of velocities centered on the systemic velocity of M31 ($v_{\text{sys}} = -300 \text{ km s}^{-1}$).

(A color version of this figure is available in the online journal.)

in (Section 4.3) that the second kinematically cold component may itself be a part of the GSS.

3.2.2. Existing GSS Core Fields at 21 kpc and 33 kpc

Fields H13s and a3 are located near the eastern edge of the GSS core at 21 and 33 kpc (Figure 1). Figures 7 and 8 display the line-of-sight velocity distributions of stars in these fields. Data in H13s and a3 have previously been presented by Kalirai et al. (2006a) and Guhathakurta et al. (2006), respectively. The data presented in this work benefit from several improvements in our reduction and analysis procedures, including an improved stellar template library for determining line-of-sight velocities, a correction to the measured velocity for mis-centering of the star in the slit, and the recovery of data from several CCD chips on the H13s_2 mask which had initially failed our automated reduction pipeline (Sections 2.2 and 2.3). Our revised estimates presented below for the mean velocity of the kinematically cold components in fields H13s and a3 are consistent with the $\sim +20 \text{ km s}^{-1}$ shift we have observed between our improved and original reduction pipelines (Figure 5).

The two kinematically cold, metal-rich components previously identified in field H13s are clearly seen in Figure 7, along with a kinematically broad, relatively metal-rich spheroid component. The parameters for the maximum-likelihood triple-

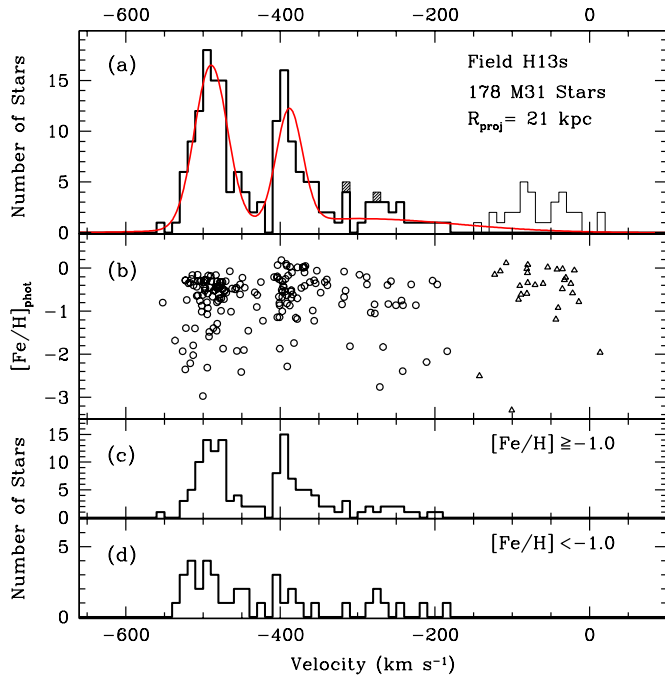


Figure 7. Line-of-sight, heliocentric velocities of stars in the GSS core field at $R_{\text{proj}} = 21$ kpc (field H13s; Section 3.2.2). Panels and symbols are the same as in Figure 6. The primary (GSS core) and the secondary kinematically cold components at $v \sim -490$ km s $^{-1}$ and $v \sim -390$ km s $^{-1}$, respectively, are clearly seen in panels (a)–(c), as well as stars related to M31’s spheroid with a wide range of velocities. The shaded boxes in panel (a) denote stars that have physical characteristics that are consistent with the M31 RGB star population, but whose very blue $(V - I)_0$ colors cause them to be classified as MW dwarf stars (Section 2.5).

(A color version of this figure is available in the online journal.)

Gaussian fit to the velocity distribution in this field are listed in Table 2.

Although both the GSS core at $v \sim -490$ km s $^{-1}$ and the second kinematically cold component ($v \sim -390$ km s $^{-1}$) in field H13s are primarily metal rich (Figures 7(b) and (c)), they are also present in the velocity distribution of stars with $[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$ (Figures 7(b) and (d)). In fact, the GSS and second kinematically cold component comprise an estimated 51% and 19% of the M31 RGB population with $[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$, respectively, similar to the respective fractions estimated using the full velocity distribution in H13s (Table 2). These percentages were estimated using a maximum-likelihood fit to the $[\text{Fe}/\text{H}]_{\text{phot}} < 1.0$ velocity distribution in Figure 7(d), with all parameters except the fraction of stars in each component held fixed at the values in Table 2.

Kalirai et al. (2006a) measured dispersions of $\sigma_v = 19$ km s $^{-1}$ for both kinematically cold components in this field, which is in agreement with our new measurements (Table 2). The median velocity measurement error is 5.4 km s $^{-1}$ in field H13s, yielding intrinsic velocity dispersions of $\sigma_v^{\text{intr}} = 20.6$ km s $^{-1}$ and $\sigma_v^{\text{intr}} = 16.1$ km s $^{-1}$ for the GSS core and second kinematically cold component, respectively.

The prominent kinematically cold, metal-rich peak in the velocity distribution of field a3, reported as a kinematical detection of the GSS by Guhathakurta et al. (2006), is evident in Figure 8 along with a significant kinematically broad spheroid component that is more metal poor, as expected for M31 RGB stars at this distance from M31 (Kalirai et al. 2006b). Table 2 lists the parameters from the maximum-likelihood double-Gaussian fit to the velocity distribution.

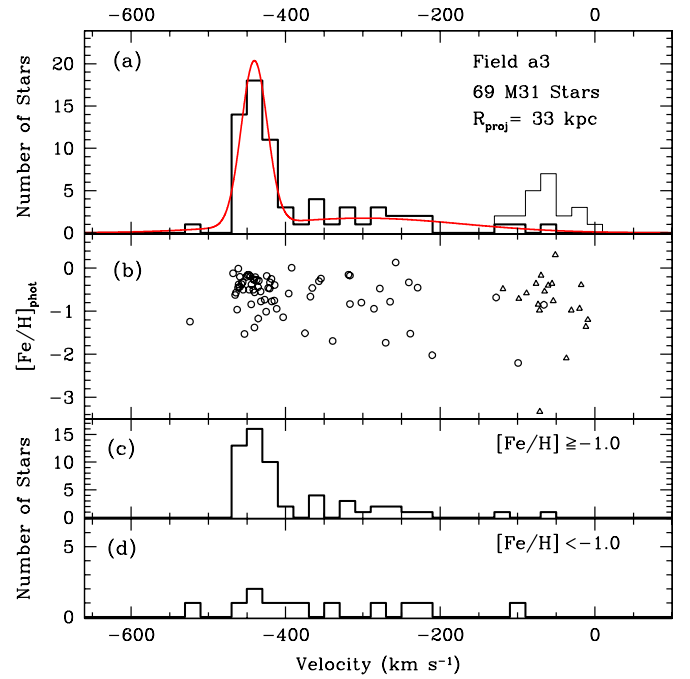


Figure 8. Line-of-sight, heliocentric velocities of stars in the GSS core field at $R_{\text{proj}} = 33$ kpc (field a3; Section 3.2.2). Panels and symbols are the same as in Figure 6. The GSS core is evident as a relatively metal-rich, kinematically cold population of RGB stars at $v \sim -440$ km s $^{-1}$ in panels (a)–(c). The velocity distribution of field a3 also has a significant contribution from M31 spheroid stars spanning a range of velocities and metallicities.

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Guhathakurta et al. (2006) measured $\sigma_v^{\text{GSS}} = 21 \pm 7$ km s $^{-1}$ for the GSS in this field, consistent with our new measurement of $\sigma_v^{\text{GSS}} = 16.8_{-3.3}^{+4.6}$ km s $^{-1}$. The median velocity measurement error in this field is 6.8 km s $^{-1}$, yielding an intrinsic velocity dispersion of $\sigma_v^{\text{intr}} = 15.4$ km s $^{-1}$ for the GSS core.

3.3. Fields Off the GSS Core

3.3.1. Fields to the West of the GSS Core

South quadrant dwarf spheroidal field at $R \sim 45$ kpc. Field d1 is located to the west of the highest-surface-brightness region of the GSS (Figure 1), yet is still solidly within the region of M31’s spheroid dominated by the GSS (as evidenced by the deeper photometric study of Ibata et al. 2007). Although it was designed to study the dwarf spheroidal galaxy And I, we have identified a total of 48 M31 field stars in addition to the large number of dSph members in this field (Section 2.5.1).

The line-of-sight velocity distribution of M31 field stars in field d1 is shown in Figure 9. There is a kinematical peak of relatively metal-rich stars at $v \sim -380$ km s $^{-1}$, and a compact group of $\sim 9\%$ of the M31 RGB stars in this field) in the v – $[\text{Fe}/\text{H}]_{\text{phot}}$ plane at $v \sim -290$ km s $^{-1}$ that are even more metal rich (Figure 9(b)). The spread in metallicity of the stars at $v \sim -290$ km s $^{-1}$ is only 0.1 dex.

M31’s halo has been shown to become increasingly metal poor toward its outskirts. The median $[\text{Fe}/\text{H}]_{\text{phot}}$ of all the M31 RGB stars in this field is -0.51 dex, significantly higher than the expected metallicity of M31’s halo at this projected distance ($[\text{Fe}/\text{H}] \lesssim -1.0$; Kalirai et al. 2006b; Chapman et al. 2006; Koch et al. 2008). The selection criterion to separate M31 field stars from dSph members only precludes stars with $[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$ and -425 km s $^{-1} < v < -375$ km s $^{-1}$ from

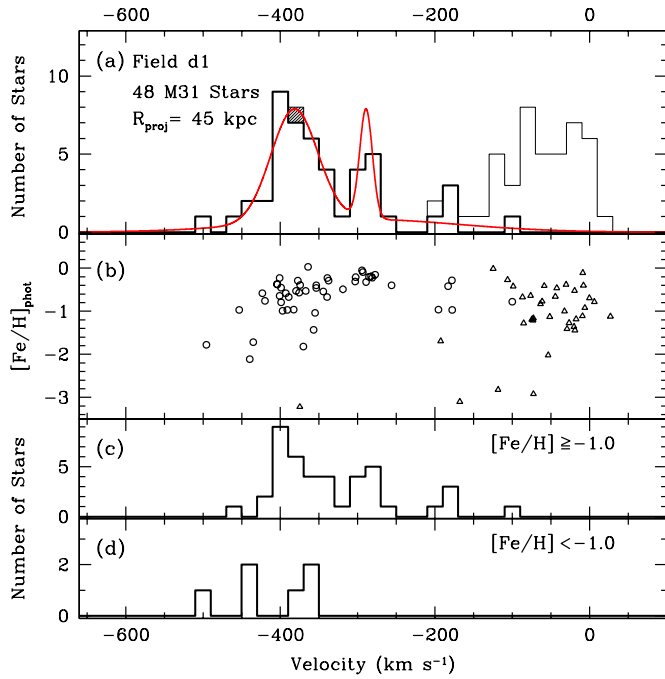


Figure 9. Line-of-sight, heliocentric velocities of stars in field d1, located in the southern quadrant of M31’s halo at $R_{\text{proj}} = 45$ kpc (Section 3.3.1). Panels and symbols are the same as in Figure 6 and the shaded boxes in panel (a) denote very blue stars as in Figure 7. Only stars that are judged unlikely to be associated with the dSph galaxy And I are shown here (Section 2.5.1). The maximum-likelihood triple-Gaussian fit (Table 2) is shown in panel (a).

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being included in the M31 RGB sample (Section 2.5.1), and thus is not a significant factor in increasing the mean metallicity of stars in the field.

The prominent kinematical peaks at $v \sim -380$ km s⁻¹ and $v \sim -290$ km s⁻¹ combined with the relatively high metallicity of the stars is strong evidence that we have detected substructure in this field. Furthermore, the data are strongly inconsistent with being drawn from a single broad ($\sigma_v \gtrsim 100$ km s⁻¹) Gaussian centered on M31’s systemic velocity (with a probability $P_{\chi^2} \ll 0.1\%$). The parameters of the maximum-likelihood triple-Gaussian fit to the M31 RGB stars in this field are listed in Table 2. A maximum-likelihood fit with the dispersion of M31’s spheroid set to the value expected for $R_{\text{proj}} \sim 45$ kpc by the Chapman et al. (2006) study (Section 3.1) returns parameters that are within the 90% confidence limits of the fit in Table 2.

The location of field d1 (Figures 1 and 5) argues strongly for the fact that the GSS is the source of at least some of the kinematic substructure seen in this field. The $v \sim -380$ km s⁻¹ kinematically cold component has the appropriate mean velocity to be the GSS (Figure 5), and we will show in Section 4.1 that the kinematical and chemical properties of the $v \sim -380$ km s⁻¹ feature are consistent with this interpretation. The nine stars with velocities ~ -290 km s⁻¹ and $[\text{Fe}/\text{H}]_{\text{phot}} \sim -0.2$ dex are evenly spread over the regions of spectroscopic masks in which we find M31 RGB field stars. They thus represent a kinematically cold tidal stream, not a spatially compact group of stars such as an intact cluster.

South quadrant M31 halo field at $R \sim 58$ kpc. Field a13 is located far to the west of the high-surface-brightness core of the GSS at $R_{\text{proj}} = 58$ kpc, solidly within the prominent “GSS envelope” photometric structure (Figure 1). The surface brightness of this field, estimated from star counts of M31

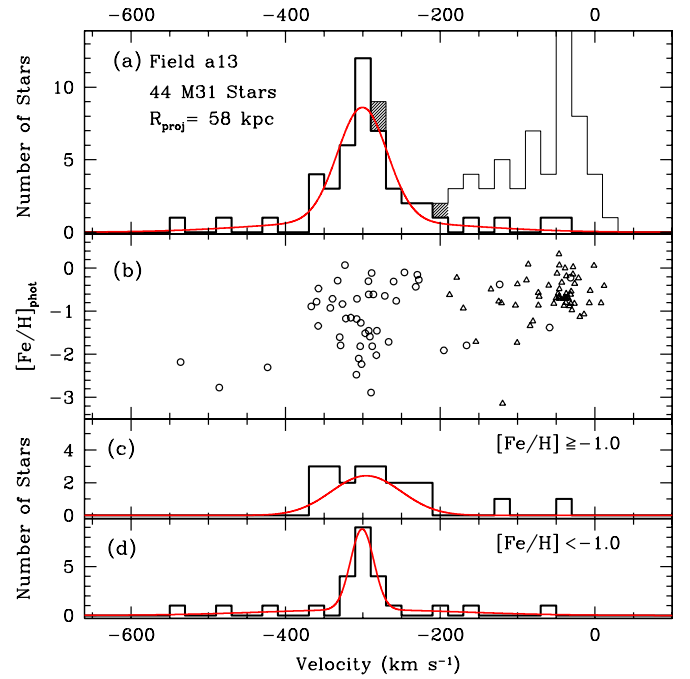


Figure 10. Line-of-sight, heliocentric velocities of stars in the south quadrant M31 halo field at $R_{\text{proj}} = 58$ kpc (field a13; Section 3.3.1) with the maximum-likelihood double-Gaussian fit (red curve) overlaid. Panels and symbols are the same as in Figure 6 and the shaded boxes in panel (a) denote very blue stars as in Figure 7. The velocity and $[\text{Fe}/\text{H}]_{\text{phot}}$ distribution of M31 RGB stars in this field shows evidence of multiple kinematical components. Relatively metal-poor M31 spheroid stars with a wide range of velocities are present, as well as a colder kinematical component centered at $v \sim -300$ km s⁻¹ (panels (b)–(d)). The relatively metal-rich subset ($[\text{Fe}/\text{H}]_{\text{phot}} \geq -1.0$, panel (c)) of the stars with $v \sim -300$ km s⁻¹ has a somewhat broader velocity dispersion than the relatively metal-poor subset ($[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$, panel (d)).

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RGB and MW dwarf stars, is significantly ($> 3\sigma$) brighter than expected based on the surface-brightness profile of M31’s stellar halo derived from star counts in fields from 18 kpc to 165 kpc (Guhathakurta et al. 2005). This implies that field a13 likely contains a significant amount of substructure which boosts its surface brightness above what is expected for a relatively smooth halo field at a radial distance of 58 kpc.

The line-of-sight velocity distribution of stars in this field is displayed in Figure 10. Stars are present with a wide range of metallicity. The relatively metal-rich stars ($[\text{Fe}/\text{H}]_{\text{phot}} \geq -1.0$) span a fairly limited range of velocities (-370 km s⁻¹ $\lesssim v \lesssim -230$ km s⁻¹), while the relatively metal-poor stars ($[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$) display both a wide range of line-of-sight velocities ($-550 \lesssim v \lesssim -150$ km s⁻¹) and a peak of stars with a narrower velocity dispersion centered at $v \approx -300$ km s⁻¹.

The parameters from the maximum-likelihood double-Gaussian fit to the full velocity distribution are listed in Table 2. If a smaller velocity dispersion of $\sigma_v = 100$ km s⁻¹ is assumed for M31’s spheroid, the resulting maximum-likelihood parameters lie well within the 90% confidence limits shown in Table 2. The kinematically cold component is estimated to comprise a large fraction ($\sim 70\%$) of the stars in this field. If the number of stars in M31’s smooth spheroid is assumed to be 30% of the M31 RGB stars in this field, the estimated spheroid surface brightness for field a13 is fully consistent with the observed surface-brightness profile of M31’s stellar halo.

The kinematical data are inconsistent ($P_{\chi^2} \ll 0.1\%$) with being drawn from a relatively broad ($\sigma_v \gtrsim 100$ km s⁻¹)

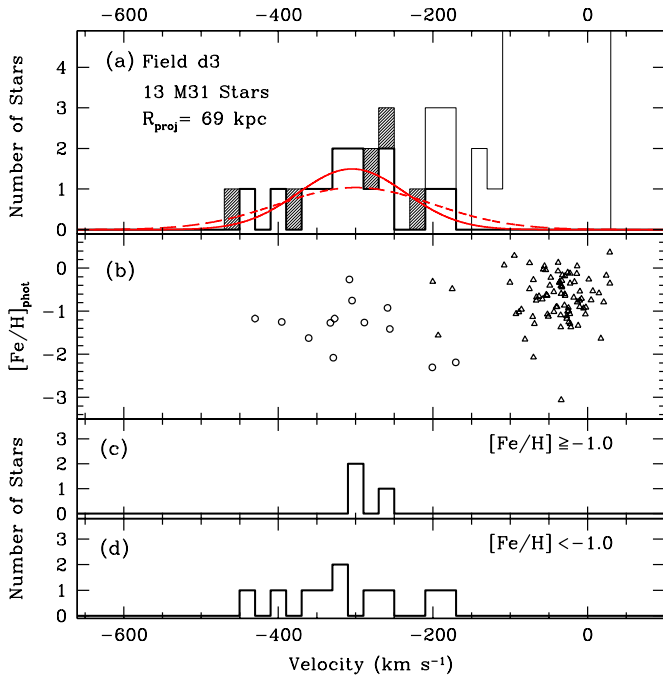


Figure 11. Line-of-sight, heliocentric velocities of stars in the dwarf spheroidal field d3, located in the southern quadrant of M31’s halo at $R_{\text{proj}} = 69$ kpc (Section 3.3.1). Panels and symbols are the same as in Figure 6 and the shaded boxes in panel (a) denote very blue stars as in Figure 7. Only stars that are unlikely to be associated with the dSph galaxy And III are shown here (Section 2.5.1). The maximum-likelihood single-Gaussian fit to the data in this field has parameters $\langle v \rangle = -305$ and $\sigma_v = 69.4$ km s $^{-1}$ (solid curve). A Gaussian with parameters $\langle v \rangle = -300$ and $\sigma_v = 100$ km s $^{-1}$ is also shown (dashed curve).

(A color version of this figure is available in the online journal.)

single Gaussian centered on M31’s velocity, but are marginally consistent with the maximum-likelihood single-Gaussian fit ($\sigma_v = 74^{+15}_{-11}$ km s $^{-1}$, $P_{\chi^2} \sim 6\%$). However, the distribution of the data in the $[\text{Fe}/\text{H}]_{\text{phot}}-v$ plane, the enhanced surface brightness observed in this field, and its location on a known photometric substructure (the GSS envelope) provide convincing evidence that there is another RGB population in this field in addition to M31’s kinematically broad spheroid. The field’s location in the photometrically prominent GSS envelope makes this the most likely source of the substructure seen in field a13.

The maximum-likelihood double-Gaussian fit to all the stars in this field is a simplification of the distribution of the data. Since one of the Gaussian components is held fixed with parameters representing M31’s kinematically broad, smooth spheroid population, only one Gaussian component is available to fit the triangular distribution seen in Figure 10(b). Panels (c) and (d) show quite clearly the difference in the velocity distributions of stars with $[\text{Fe}/\text{H}]_{\text{phot}} \geq -1.0$ and $[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$ dex. Gaussian fits to the velocity distribution of the stars in each of these panels have been overlaid. We assume a broad kinematical component is present in the $[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$ population, but not in the $[\text{Fe}/\text{H}]_{\text{phot}} \geq -1.0$ population. This is motivated in part by the velocity distribution of the stars (there are clearly metal-poor stars spread over a very large range in velocity in Figure 10(d) while a similar distribution is not seen in Figure 10(c)) and in part by the fact that M31’s stellar halo has been shown to be largely metal poor at these large radii ($\langle [\text{Fe}/\text{H}]_{\text{phot}} \rangle \sim -1.5$; Kalirai et al. 2006b; Koch et al. 2008). The velocity dispersion of the kinematical peak of stars with $[\text{Fe}/\text{H}]_{\text{phot}} \geq -1.0$ is $\sigma_v = 45^{+15}_{-10}$ km s $^{-1}$, and

the velocity dispersion of the kinematical peak of stars with $[\text{Fe}/\text{H}]_{\text{phot}} < -1.0$ is $\sigma_v = 15.1^{+7.4}_{-5.2}$ km s $^{-1}$. These values are both marginally consistent with the dispersion found in the fit to the full velocity distribution (Table 2). The mean velocities are fully consistent with each other and with the results of the fit to the full velocity distribution.

We can only speculate for now on the relationship between the relatively metal-poor, narrower kinematical peak seen in Figure 10(d) and the relatively metal rich, slightly broader kinematical peak in Figure 10(c). Another pointing in a different region of the GSS envelope is necessary to determine if this structure in $[\text{Fe}/\text{H}]-v_{\text{los}}$ space persists throughout the GSS envelope region, or is specific to the location of field a13. If it persists, the interesting structure seen in the $[\text{Fe}/\text{H}]_{\text{phot}}-v$ is likely endemic to the stars associated with the GSS envelope and represents one coherent substructure. If only part of this structure is seen in another location on the GSS envelope, (i.e., only the relatively metal-poor, narrower kinematical peak or only the relatively metal rich, slightly broader kinematical peak), it would give credence to the interpretation of this structure as two independent kinematical substructures in M31’s halo which are superimposed along our line of sight. For the present, we use the fit to the full sample of stars in this field (Table 2), and treat the substructure as one (complex) kinematical component, identified as the first kinematical detection of the GSS envelope.

South quadrant dwarf spheroidal field at $R \sim 69$ kpc. Field d3 is far to the west of the high-surface-brightness GSS core, and on the edge of the region of M31’s stellar halo which appears contaminated by the GSS in star-count maps (Figure 1). Although the spectroscopic masks were designed to study RGB members of the dwarf spheroidal galaxy And III, we have also identified a total of 13 M31 field stars (Section 2.5.1).

The line-of-sight velocity distribution of the 13 M31 field stars in field d3 is shown in Figure 11. The M31 RGB stars span a wide range in velocity and $[\text{Fe}/\text{H}]_{\text{phot}}$ values, and the velocities are consistent with being drawn from either a Gaussian with parameters $\langle v \rangle^{\text{sph}} = -300$ km s $^{-1}$, $\sigma_v^{\text{sph}} = 129$ km s $^{-1}$ or $\sigma_v^{\text{sph}} = 100$ km s $^{-1}$ (the dispersion implied for the projected radius of this field by the Chapman et al. (2006) study). The maximum-likelihood single Gaussian in this field has parameters $\langle v \rangle = -304.8^{+33.3}_{-33.8}$ and $\sigma_v = 69.4^{+30.2}_{-17.5}$ km s $^{-1}$. Based on the current limited sample of stars, there is no kinematical evidence of substructure in this field.

3.3.2. Southeast Minor-axis Field at $R \sim 57$ kpc

Field m4 is located on M31’s southeast minor axis at $R_{\text{proj}} \sim 57$ kpc (Figure 1), at the same location as the arc-like photometric substructure termed “stream C.” The line-of-sight velocity distribution of stars in this field is displayed in Figure 12. Two prominent kinematically cold features are present: a relatively metal-rich ($[\text{Fe}/\text{H}]_{\text{phot}} \gtrsim -1.0$) group of stars at $v \sim -350$ km s $^{-1}$ and a secondary group of stars with somewhat lower metallicities at $v \sim -260$ km s $^{-1}$. Another spectroscopic study of M31 RGB stars $\sim 1^\circ$ southwest of field m4 also found evidence of two kinematically cold components (Chapman et al. 2008; the mean velocities and dispersions of these components are designated by the arrows in Figure 12). In Section 4.2, we will show that both kinematically cold components can be identified with photometric substructure seen in the region of stream C.

The parameters of the maximum-likelihood triple-Gaussian fit to stars in field m4 are presented in Table 2. The maximum-likelihood triple-Gaussian parameters assuming a smaller

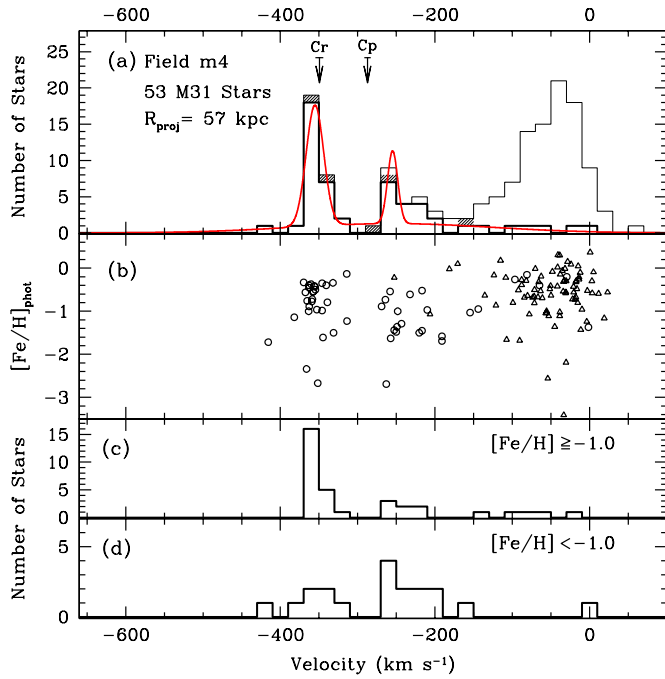


Figure 12. Line-of-sight, heliocentric velocities of stars in the southeast minor-axis field at $R_{\text{proj}} = 57$ kpc (field m4; Section 3.3.2). Panels and symbols are the same as in Figure 6. There is a prominent, kinematically cold peak of metal-rich stars in this field centered at $v \sim -360$ km s $^{-1}$, and a second kinematically cold peak of more metal-poor stars at $v \sim -260$ km s $^{-1}$. The arrows denote the mean velocities and dispersions of kinematically cold features detected in a nearby spectroscopic field (Chapman et al. 2008). As in Figure 7, shaded boxes in panel (a) denote blue stars for which our M31 RGB/MW dwarf classification is less certain (Section 2.5).

(A color version of this figure is available in the online journal.)

velocity dispersion for M31’s spheroid ($\sigma_v = 100$ km s $^{-1}$) are well within the 90% confidence limits in Table 2. The median velocity measurement error in field m4 is 5.4 km s $^{-1}$. Thus, the intrinsic velocity dispersions of the primary and secondary kinematically cold components are $\sigma_v^{\text{intr}} = 10.0$ km s $^{-1}$ and 3.8 km s $^{-1}$, respectively; the secondary kinematically cold component is just resolved.

3.4. Summary of the Results for Individual Fields

We have detected and measured the mean velocity and velocity dispersion of the GSS in 5 fields: f207, H13s, a3, d1, and a13. Of these, field f207 provides the innermost measurement of the mean velocity and velocity dispersion of the GSS at $R_{\text{proj}} \sim 17$ kpc, and field a13 provides the first measurement of the kinematical properties of the GSS envelope. Four of our seven fields show not one but *two* kinematically cold components in addition to the kinematically hot M31 spheroid population. Field f207 has a second kinematically cold component with a mean velocity ~ 100 km s $^{-1}$ more positive than that of the GSS (Section 3.2.1). This component mirrors the secondary kinematically cold component in field H13s, which also has a mean velocity ~ 100 km s $^{-1}$ more positive than the GSS’s. The second kinematically cold components in fields f207 and H13s likely have the same physical origin as each other, and may also be part of the GSS (Section 4.3). The velocity distribution of field m4 shows two prominent, kinematically cold peaks (Section 3.3.2). They confirm a previous finding of two kinematically cold components in the photometrically identified minor-axis arc stream C.

4. A COMPARISON OF THE KINEMATICAL AND CHEMICAL ABUNDANCE PROPERTIES OF THE GSS AND OTHER DEBRIS

In Section 3, we identified and characterized kinematically cold components in individual fields. In this section, we widen our focus and discuss the chemical and kinematical properties of these components in relation to each other and in the context of physical structures known to exist in M31’s spheroid. We also examine the evidence for various physical origins for the kinematical components which are not clearly attributable to the GSS: the kinematically cold components in field m4 (Section 4.2) and the second kinematically cold components in fields f207 and H13s (Section 4.3.1).

4.1. GSS Core and Envelope

Debris from the GSS extends over a large portion of the southern-most quadrant of M31’s stellar halo. The structure of the GSS includes the high-surface-brightness, high metallicity “core” ($\langle[\text{Fe}/\text{H}]\rangle = -0.54$) surrounded by a more diffuse, lower metallicity “envelope” ($\langle[\text{Fe}/\text{H}]\rangle = -0.71$; Ibata et al. 2007). Numerical simulations of the collision of the GSS’s progenitor with M31 show that the core/envelope structure is naturally produced if the progenitor is a disk galaxy with an internal metallicity gradient (Fardal et al. 2008).

We have presented kinematical detections of the GSS core in fields f207 (Section 3.2.1), and H13s and a3 (Section 3.2.2). We have also detected kinematically cold components in fields d1 and a13 which we attribute to the GSS core and GSS envelope (Section 3.3.1), respectively. We will show below that their respective kinematical and chemical properties corroborate this interpretation.

4.1.1. Kinematics

An inspection of the velocity distributions of fields on the GSS clearly shows that its velocity (marked with arrows in Figure 5) becomes increasingly positive with increasing distance from the center of M31 ($\langle v \rangle_{\text{GSS}} = -527$ km s $^{-1}$ in field f207, while $\langle v \rangle_{\text{GSS}} = -318$ km s $^{-1}$ in field IF1). Figure 13 displays the mean line-of-sight velocity and observed velocity dispersion (Table 2) of the GSS as a function of projected radius from M31’s center. The mean velocity and velocity dispersion of the GSS core in the Ibata et al. (2004) fields (Figure 5) were calculated using maximum-likelihood double-Gaussian fits, with $G^{\text{sph}}(v)$ held fixed with parameters $\langle v \rangle = -300$ km s $^{-1}$ and $\sigma_v^{\text{sph}} = 129$ km s $^{-1}$ (Section 3.1). The error bars represent the 90% confidence limits from the maximum-likelihood analysis.

We performed a weighted least-squares linear fit to the mean line-of-sight velocity of the GSS as a function of R_{proj} . The result, using the fields on the highest-surface-brightness region of the GSS core, is

$$\langle v \rangle_{\text{GSS}} = [63.0(R_{\text{proj}}/1^\circ) - 594.4] \text{ km s}^{-1}. \quad (1)$$

This fit includes the data from fields f207, H13s, IF6, a3, IF2, and IF1 (Figure 1). These fields are represented by open circles and triangles in Figure 13. The open circle at 3 $^{\circ}$ 25, field d1, was not included in the above fit and is discussed further below.

Although the line-of-sight velocity of the orbit of the GSS progenitor is expected to be strongly nonlinear over the range of radii observed, the GSS exists because the progenitor and stars belonging to the GSS are on slightly different orbits. When this is taken into account in modeling the interaction which produced

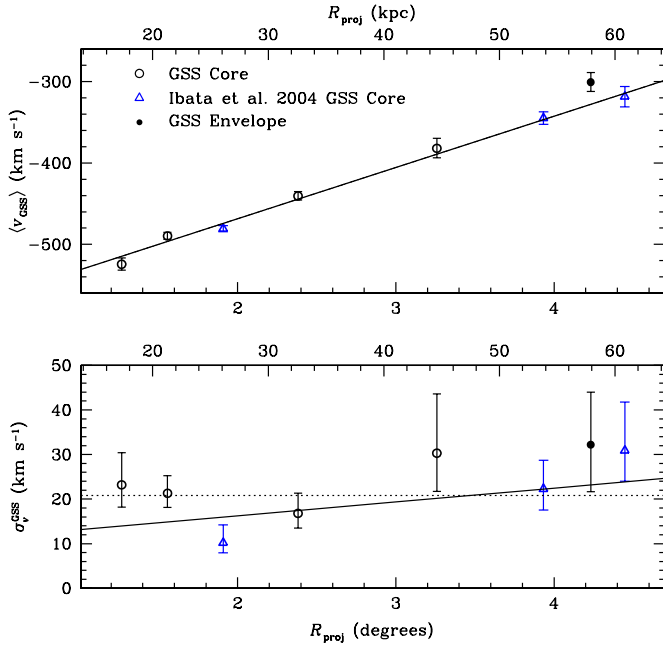


Figure 13. Mean line-of-sight velocity and observed line-of-sight velocity dispersion (not corrected for measurement error) of stars associated with the GSS as a function of radial projected distance from the center of M31 (Section 4.1.1). The values are from maximum-likelihood, multiple-component Gaussian fits (Section 3) to the stellar velocity distributions of each of the fields, including GSS core fields (open circles), the GSS envelope field (filled circle), and Ibata et al. (2004) GSS core fields (open triangles). Error bars denote 90% confidence limits from the maximum-likelihood analysis. Weighted least-squares linear fits to the GSS core data from fields f207, H13s, IF6, a3, IF2, and IF1 are also shown. The dotted line in the bottom panel shows the mean velocity dispersion of these fields. The mean line-of-sight velocity and velocity dispersion of the GSS in fields d1 and a13 (open circle at 3°25 and solid circle, respectively) are consistent with the fits to the core fields along the eastern edge of the GSS.

(A color version of this figure is available in the online journal.)

the GSS, the orbital models produce a roughly linear trend for the mean line-of-sight velocity of the GSS as a function of radial distance over most of the observed distance range (Fardal et al. 2006). The linear approximation is an adequate description of the data over the range of radii for which the velocity of the GSS has been measured.

The d1 field ($R_{\text{proj}} \sim 45$ kpc) is very near the region chosen by Ibata et al. (2007) for analysis of the GSS core's stellar population and is consistent with an extrapolation of this region toward M31's center (Figure 1). The GSS core velocity at the projected radius of our d1 field, based on the above linear fit, is expected to be $\langle v \rangle \sim -390$ km s $^{-1}$. This is consistent with the velocity found for the primary kinematically cold component in field d1 (Table 2, Section 3.3.1).

In addition to the fields in the GSS core, Figure 13 also displays the mean line-of-sight velocity and velocity dispersion of the kinematically cold component in the a13 field (filled circle). The mean velocity of the kinematically cold component in field a13 is more positive than expected from the linear fit to the core fields. The mean velocity of the GSS envelope is expected to be relatively close to, but not necessarily the same, as the mean velocity of the GSS core at the same projected radius (M. Fardal 2009, private communication).

There is not a significant trend with R_{proj} in the line-of-sight velocity dispersion of stars in the GSS (Figure 13, bottom panel), although the data hint that the velocity dispersion in the outer regions of the GSS ($R_{\text{proj}} > 40$ kpc) may be slightly larger

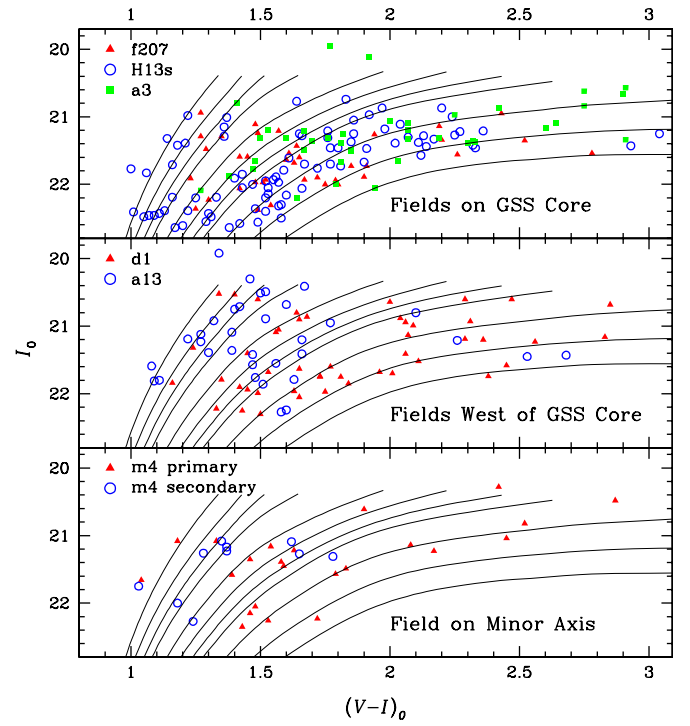


Figure 14. Position of stars in the $[I_0, (V - I)_0]$ CMD for stars within $\pm 2\sigma_v$ of the mean velocity of the primary substructure in fields on and off the GSS core (Section 4.1.2). Isochrones shown are 12 Gyr, $[\alpha/\text{Fe}] = 0.0$ from Vandenberg et al. (2006) and are the same set shown in Figure 3. The most metal-poor (bluest) isochrone represents $[\text{Fe}/\text{H}] = -2.31$ and the most metal-rich (reddest) isochrone represents $[\text{Fe}/\text{H}] = 0.0$. The CMD distribution of GSS stars in the three core GSS fields is similar, while the stars in the substructure in a13 and the secondary kinematically cold component in m4 are on average more blue. (A color version of this figure is available in the online journal.)

on average than the inner regions ($R_{\text{proj}} < 40$ kpc). The mean velocity dispersion of the GSS in the fields along its eastern edge is $\langle \sigma_v \rangle = 20.8$ km s $^{-1}$. The mean value is just outside the 90% confidence limits of the velocity dispersion measured in the more western core field d1 and the GSS envelope field a13. The weighted least-squares fit to the eastern GSS core fields is also shown in Figure 13. The velocity dispersion of the GSS in fields d1 and a13 is consistent with this linear fit.

4.1.2. Metallicity Distributions

Figure 14 shows the position in an $[I_0, (V - I)_0]$ CMD of all stars within $\pm 2\sigma_v$ of the mean velocity of the primary kinematical substructure in fields on and off the GSS core. The stars within $\pm 2\sigma_v$ of the GSS core in fields f207, H13s, and a3 have similar distributions in color-magnitude space. The stars within $\pm 2\sigma_v$ of the kinematically cold component in field a13, located on the GSS envelope, are significantly bluer than the stars associated with the GSS core. We have not corrected for the varying line-of-sight distance to the GSS in Figure 14. The difference in the line-of-sight distance to the GSS between fields a3 and f207 is ~ 20 kpc (McConnachie et al. 2003), which corresponds to only a 0.05 dex difference in apparent magnitude.

Differences in the CMD distributions of RGB stars are caused by differences in the metallicity, age, or alpha enhancement of the stellar populations. For sufficiently old stellar populations (older than ~ 4 Gyr; Gallart et al. 2005), the RGB is primarily sensitive to metallicity, with age and alpha enrichment playing secondary roles. Although the alpha enrichment of RGB stars in

M31's spheroid and the GSS is not observationally constrained, these stars are known to span a range of ages (Section 2.4), in contrast to the 12 Gyr old coeval population we assume in measuring $[\text{Fe}/\text{H}]_{\text{phot}}$. With these caveats, we interpret differences in the CMD distribution of stars as differences in metallicity in the following analysis. Due to the age–metallicity degeneracy of the RGB, the *differences* between the metallicity distributions of features, which reflect the observed differences in the CMD distributions, deserve more weight than the absolute metallicity estimates of the stars.

The assumed distance to the stellar population also affects the photometric metallicity estimates of stars. However, the line-of-sight distance to the GSS core in individual fields is not known with a great deal of certainty, and the line-of-sight distances to the other kinematic features identified in Section 3 are not yet known. We therefore assume M31's distance modulus when deriving the photometric metallicity estimates. For the GSS kinematic features, we will note the effect of assuming the distance measured by McConnachie et al. (2003) on the $[\text{Fe}/\text{H}]_{\text{phot}}$ value.

The final source of systematic error in deriving the mean metallicity abundance of a kinematic feature is the assignment of individual stars to kinematic components. Individual stars can be associated with the kinematically cold components only statistically (i.e., through maximum-likelihood Gaussian fits as in Section 3). Therefore, the $[\text{Fe}/\text{H}]$ distributions discussed below include stars associated with the kinematically cold components as well as stars associated with M31's spheroid that happen to have velocities within $\pm 2\sigma_v$ of the mean velocity of the kinematically cold components. For the GSS core fields, this is a relatively small effect since the GSS core is a narrow and dominant feature well offset from M31's systemic velocity (Figure 5).

Figure 15 displays the $[\text{Fe}/\text{H}]$ distributions of the stars in Figure 14 in differential and cumulative form. The GSS core has a very similar $[\text{Fe}/\text{H}]$ distribution in all three GSS core fields (Figure 15(a)), with a strong metal-rich peak and an extended metal-poor tail. The median metallicity of stars within $\pm 2\sigma_v$ of the mean velocity of the GSS component is $\langle [\text{Fe}/\text{H}] \rangle_{\text{med}} = -0.55$, -0.51 , and -0.44 in fields f207, H13s, and a3, respectively; $\langle [\text{Fe}/\text{H}] \rangle_{\text{med}} = -0.49$ ($\langle [\text{Fe}/\text{H}] \rangle_{\text{mean}} = -0.66 \pm 0.04$) for all three fields combined.

Accounting for the estimated distance to the GSS in each of the above fields (McConnachie et al. 2003) results in only a ~ -0.04 dex difference in the derived $[\text{Fe}/\text{H}]_{\text{phot}}$ estimates, while changing the assumption of the age of the population from 12 Gyr to 8 Gyr results in a $\sim +0.1$ dex difference in the derived $[\text{Fe}/\text{H}]_{\text{phot}}$ estimates. Based on the multi-Gaussian fits, between 5% and 8% (or ~ 3 –4 stars) of the red giants within $\pm 2\sigma_v$ of the mean velocity of the GSS core in fields f207, H13s, and a3 are associated with M31's spheroid. Within $R_{\text{proj}} \sim 30$ kpc, M31's spheroid is found to be relatively metal rich (Kalirai et al. 2006b; Gilbert et al. 2007; Koch et al. 2008). However, if we assume that the GSS core is more metal rich than M31's spheroid, and remove the 3–4 most metal-poor stars in each field, the resulting median $[\text{Fe}/\text{H}]_{\text{phot}}$ estimate for the GSS increases by only 0.02 dex.

The $[\text{Fe}/\text{H}]$ distribution of stars within $\pm 2\sigma$ of the mean velocity of the -380 km s^{-1} kinematically cold component in field d1 is very similar to the $[\text{Fe}/\text{H}]$ distribution of stars associated with the GSS core (Figure 15(b)). The median metallicity of the stars associated with the GSS in field d1 is ~ 0.1 dex lower than the median

$[\text{Fe}/\text{H}]$ of the GSS core in fields f207, H13s, and a3. If the estimated distance to the GSS in each of the fields is accounted for, the median metallicity of the GSS in field d1 is found to be 0.13 dex more metal poor than in the more eastern GSS core fields f207, H13s, and a3.

Stars within $\pm 2\sigma_v$ of the mean velocity of the kinematically cold component in field a13 are on average significantly more metal poor than the GSS core (Figure 15(b)). The median metallicity of the stars associated with the kinematically cold component in field a13 is ~ 0.7 dex lower than the median $[\text{Fe}/\text{H}]$ of the GSS core ($\langle [\text{Fe}/\text{H}] \rangle_{\text{med}} = -1.18$, $\langle [\text{Fe}/\text{H}] \rangle_{\text{mean}} = -1.19 \pm 0.12$). This is qualitatively consistent with results from purely photometric data that the GSS envelope is more metal poor than the GSS core (Ibata et al. 2007). A two-sided Kolmogorov–Smirnov (KS) test returns a probability of 0.001% that the metallicity distribution function (MDF) of the kinematically cold component in field a13 and the MDF of the GSS core were drawn from the same distribution. If we assume that the GSS envelope follows the same trend of increasing line-of-sight distance with R_{proj} as the GSS core and correct for this effect, the difference in the metallicity between the GSS core in fields f207, H13s, and a3 and the GSS envelope in field a13 becomes slightly greater (~ 0.8 dex). The largest systematic effect in this field may come from the identification of individual stars as part of the kinematically cold component. The double-Gaussian fit to the stars in this field (Table 2) implies that 13.5% of the stars (~ 5 stars) within $\pm 2\sigma_v$ of the mean velocity of the kinematically cold component are associated with the spheroid. If we assume that the five lowest metallicity stars within $\pm 2\sigma_v$ of the kinematically cold component are M31 halo stars, the estimated $[\text{Fe}/\text{H}]_{\text{phot}}$ of the GSS envelope debris increases by $+0.25$ dex ($\langle [\text{Fe}/\text{H}] \rangle_{\text{med}} = -0.92$, $\langle [\text{Fe}/\text{H}] \rangle_{\text{mean}} = -1.00 \pm 0.10$).

4.2. Minor-axis Arc at $R_{\text{proj}} \sim 57 \text{ kpc}$

We identified two kinematically cold components in field m4: one at $v = -355 \text{ km s}^{-1}$ which is relatively metal rich and the other at $v = -255 \text{ km s}^{-1}$ which is more metal poor (Section 3.3.2; Figure 12). This field is located squarely on the 60 kpc minor-axis arc “stream C” (Figure 1), identified as a photometric overdensity in star-count maps. The M31 star-count map does show evidence of two overdensities in the region of stream C: a more prominent overdensity that is relatively metal rich and a less prominent overdensity that is relatively metal poor (Chapman et al. 2008). Both the relative contributions of the kinematically cold components in field m4 (Table 2) and their metallicities (Figure 12, discussed in detail below) imply that the $v \sim -355 \text{ km s}^{-1}$ component corresponds to the former, and the $v \sim -255 \text{ km s}^{-1}$ component corresponds to the latter.

Our field is 1:1 northeast of the stream C spectroscopic fields F25 and F26 presented by Chapman et al. (2008). They also find two kinematically cold components in this region: a relatively metal-rich stream “Cr” ($\langle v \rangle = -349.5 \text{ km s}^{-1}$ and $\sigma_v^{\text{intr}} = 5.1 \pm 2.5 \text{ km s}^{-1}$ based on nine stars (quoted errors are 1σ)) and a relatively metal-poor stream “Cp” ($\langle v \rangle = -285.6 \text{ km s}^{-1}$ and $\sigma_v^{\text{intr}} = 4.3^{+1.7}_{-1.4} \text{ km s}^{-1}$ based on six stars). Arrows in Figure 12(a) mark the mean line-of-sight velocities and intrinsic velocity dispersions of streams Cr and Cp in the F25/F26 field. Due to the spatial proximity of our m4 spectroscopic masks and the Chapman et al. (2008) F25/F26 masks as well as the proximity of the kinematical features in velocity space, it is reasonable to identify our $v = -355 \text{ km s}^{-1}$ component with stream “Cr” and our $v = -255 \text{ km s}^{-1}$ component with stream

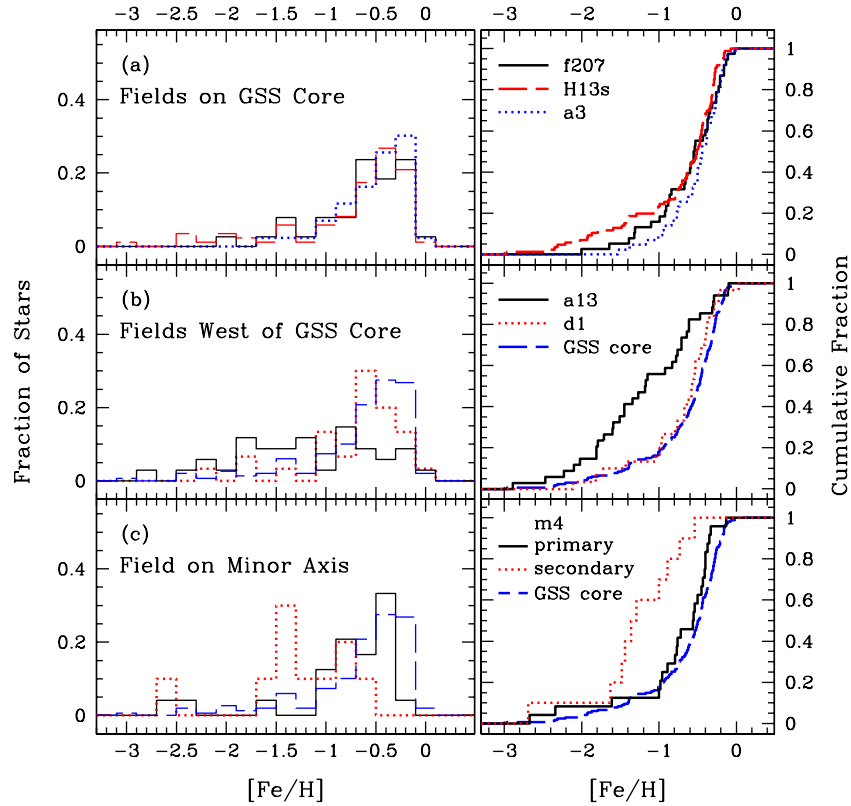


Figure 15. Metallicity distribution function (MDF) of stars within $\pm 2\sigma_v$ of the mean velocity of the primary substructure in fields on and off the GSS core, in histogram (left) and cumulative (right) form. (a) MDF of the GSS core components in fields f207 (solid histogram and curve), H13s (dashed histogram and curve), and a3 (dotted histogram and curve), discussed in Section 4.1.2. (b) MDF of stars associated with the kinematically cold component in the d1 (dotted histogram and curve) and a13 (solid histogram and curve) fields. The combined MDF of the GSS core components in panel (a) is shown for comparison (dashed histogram and curve, Section 4.1.2). (c) MDF of stars associated with the $v \sim -355 \text{ km s}^{-1}$ (solid histogram and curve) and $v \sim -255 \text{ km s}^{-1}$ (dotted histogram and curve) kinematically cold components in field m4 (Section 4.2). As in panel (b), the MDF of the combined GSS core components is shown for comparison (dashed histogram and curve). The GSS core MDFs from fields f207, H13s, and a3 are consistent with one other and are similar to those of the d1 field and the m4 primary kinematically cold component, while the MDFs of the GSS envelope in field a13 and the secondary kinematically cold component in field m4 are significantly more metal poor.

(A color version of this figure is available in the online journal.)

“Cp.” A comparison of the mean velocities of the kinematically cold components in the m4 and F25/F26 fields suggests that their velocities diverge as one moves from the southwest to the NE along stream C.

The $[I_0, (V - I)_0]$ CMD distributions of stars within $\pm 2\sigma_v$ of the mean velocities of the $v = -355 \text{ km s}^{-1}$ and $v = -255 \text{ km s}^{-1}$ kinematically cold components in field m4 are shown in the bottom panel of Figure 14. It is clear that on average the stars in the $v = -355 \text{ km s}^{-1}$ primary component are redder than the stars in the $v = -255 \text{ km s}^{-1}$ secondary component. The $[\text{Fe}/\text{H}]$ distributions of these stars (assuming a coeval population of 12 Gyr) is displayed in Figure 15(c). The differences seen in the $(I, V - I)$ CMD translate to clear differences in the MDFs of the stars. While the $[\text{Fe}/\text{H}]$ distribution of the $v = -355 \text{ km s}^{-1}$ primary cold component in field m4 is relatively metal rich and similar to the $[\text{Fe}/\text{H}]$ distribution of the GSS, the MDF of the $v = -255 \text{ km s}^{-1}$ secondary cold component is significantly more metal poor.

The median metallicity of the $v \sim -355 \text{ km s}^{-1}$ kinematically cold component is $\langle [\text{Fe}/\text{H}] \rangle_{\text{med}} = -0.56$ ($\langle [\text{Fe}/\text{H}] \rangle_{\text{mean}} = -0.79 \pm 0.12$), which is similar to the metallicity measured for stream C of $[\text{Fe}/\text{H}] = -0.6$ (Ibata et al. 2007; measured from the region shown in Figure 1). Using 10 Gyr old Padova isochrones (Girardi et al. 2002), Chapman et al. (2008) find an average metallicity of $\langle [\text{Fe}/\text{H}] \rangle_{\text{avg}} = -0.74 \pm 0.19$ for stream Cr.

Based on the double-Gaussian fit to the velocities in this field (Table 2), 14% of the stars within $\pm 2\sigma_v$ of the $v \sim -355 \text{ km s}^{-1}$ kinematically cold component in field m4 are likely associated with M31’s spheroid (~ 3 stars). M31’s spheroid has been found to be metal poor at large radii (Kalirai et al. 2006b; Koch et al. 2008). If we assume that the lowest metallicity stars are those associated with M31’s spheroid, the estimated median $[\text{Fe}/\text{H}]_{\text{phot}}$ of the kinematically cold component increases by only $+0.02$ dex.

The $v \sim -255 \text{ km s}^{-1}$ kinematically cold component is on average more metal poor than the $v \sim -355 \text{ km s}^{-1}$ component (Figure 15(c)). The median metallicity of stars within $\pm 2\sigma_v$ of the mean velocity is $\langle [\text{Fe}/\text{H}] \rangle_{\text{med}} = -1.33$ ($\langle [\text{Fe}/\text{H}] \rangle_{\text{mean}} = -1.31 \pm 0.18$). Using 10 Gyr old Padova isochrones, Chapman et al. (2008) find an average metallicity of $\langle [\text{Fe}/\text{H}] \rangle_{\text{mean}} = -1.26 \pm 0.16$ for stream Cp.

A physical association with the GSS? Based on an analysis of all stars within the stream C and GSS core regions marked in Figure 1, Ibata et al. (2007) claimed an association between the GSS and the 60 kpc minor-axis arc was unlikely due to significant differences in the properties of their stellar populations. However, if the GSS’s progenitor was a disk galaxy with a reasonable metallicity gradient, simulations show that arc-like features along M31’s southeast minor axis with different stellar populations from the GSS can be produced during the progenitor’s disruption (Fardal et al. 2008), although

they are farther from M31's center than the observed minor-axis arcs. The simulated arc features have mean line-of-sight velocities that are close to the systemic velocity of M31 ($v_{\text{sys}} = -300 \text{ km s}^{-1}$). There is also observational precedence for variations in metallicity along a tidal feature: a strong metallicity gradient has been observed in the MW's Sagittarius stream (Chou et al. 2007). Therefore, it remains possible that the GSS and some of the substructure seen in this region share a common progenitor. Regardless, the relatively high metallicity of the $v = -355 \text{ km s}^{-1}$, “Cr” stream indicates its progenitor was likely fairly massive (e.g., Larson 1974; Mateo 1998; Dekel & Woo 2003).

Connection with the extended cluster EC4. Collins et al. (2009) find that stream Cp from the Chapman et al. (2008) study has a line-of-sight velocity and metallicity consistent with the extended cluster EC4, which was discovered as part of the CFHT/Megacam survey of M31 (Ibata et al. 2007) and had RGB members which were targeted on the F25/F26 masks. EC4 has a half-light radius of 33.7 pc and $M_V = 6.6$ (Mackey et al. 2006), placing it solidly in the gulf between globular clusters and dwarf spheroidals in size–luminosity space. One proposed explanation for such objects is that they are the stripped cores of dSphs. It is thus tempting to consider the possibility that stream Cp is the tidal stream of a stripped dSph, whose core remains intact as the cluster EC4. However, the measured M/L ratio for EC4 of $6.7^{+15}_{-6.7} M_{\odot}/L_{\odot}$ (Collins et al. 2009) makes this unlikely. Collins et al. (2009) argue that it is also unlikely the stream Cp stars are tidally disrupted members of EC4 itself, since they are between 15 and 74 core radii distant from EC4 and *HST* photometry of EC4 shows no evidence of distorted isophotes or significant tidal debris (Mackey et al. 2006). Our discovery of the continuation of stream Cp in a field 1:1 (15 kpc) to the northeast confirms that stream Cp is not formed from tidally stripped cluster stars. It is most likely that stream Cp and the $v = -255 \text{ km s}^{-1}$ stream in field m4 represent the tidal stream of an unknown progenitor, from which the extended cluster EC4 was also tidally stripped.

In conclusion, we find that the metallicity distribution of stars kinematically associated with the $v \sim -355 \text{ km s}^{-1}$ cold component and the GSS core are similar (Figure 15), while the metallicity distribution of stars associated with the $v \sim -255 \text{ km s}^{-1}$ is significantly more metal poor than the GSS core. The two kinematically cold components in field m4 are consistent with being detections further along stream C of the kinematical substructures “Cr” and “Cp” (Chapman et al. 2008). The progenitors of these streams remain unknown.

4.3. Second Kinematically Cold Component in the GSS Core Fields

4.3.1. Radial Trends

Each of the two innermost GSS core fields (adjacent to each other at $R_{\text{proj}} \sim 17$ and 21 kpc) has a second kinematically cold component in addition to the primary kinematical component identified as the GSS core (Figure 5, Sections 3.2.1 and 3.2.2). The top panel of Figure 16 displays line-of-sight velocity versus R_{proj} for the M31 RGB stars in fields f207 and H13s. Stars within $\pm 2\sigma_v$ of the mean velocity of the GSS core and second kinematically cold component in each field (Table 2) are denoted by blue and red points, respectively. Stars comprising the main GSS peak and the second cold component follow nearly parallel linear trends in the line-of-sight velocity versus R_{proj} plane. These trends are continuous across the boundary between the two fields.

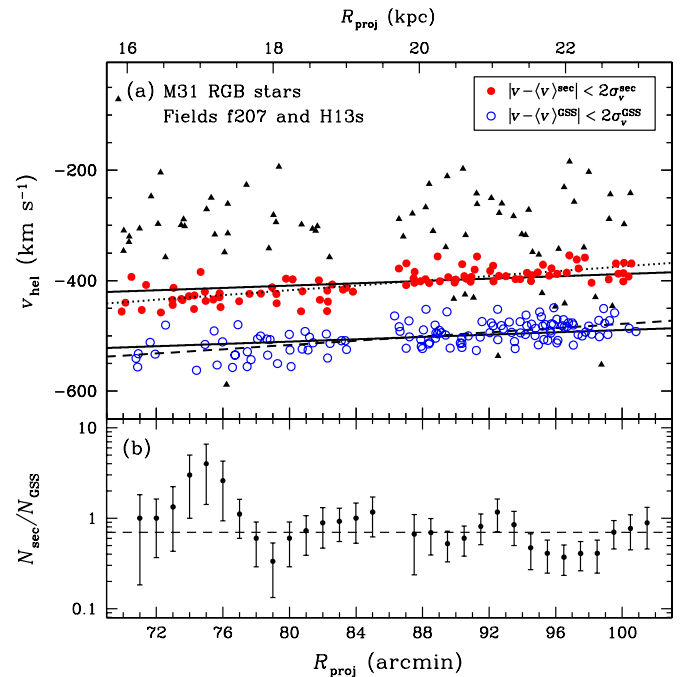


Figure 16. Comparison of stars associated with the GSS core and the second kinematically cold component in fields f207 and H13s, at $R_{\text{proj}} \sim 17$ and 21 kpc, respectively (Section 4.3.1). (a) The line-of-sight velocity distribution of M31 RGB stars in fields f207 and H13s as a function of projected radial distance from M31's center. Open circles denote stars within $\pm 2\sigma_v$ of the mean velocity of the GSS core in each field, while solid circles denote stars within $\pm 2\sigma_v$ of the mean velocity of the second kinematically cold component in each field. Solid triangles denote the remainder of the M31 RGB stars in these fields. The solid line is the linear fit to the mean velocity of the GSS core as a function of radius (Figure 13, Section 4.1.1), offset by +100 km s⁻¹ for the second kinematically cold component. The dashed and dotted lines are the linear least-squares fits to the GSS core and second kinematically cold component data, respectively. (b) The ratio of the number of stars in the second kinematically cold component and the GSS core ($N_{\text{sec}}/N_{\text{GSS}}$) as a function of radius. The ratio was calculated using running bins of width 3'. The error bars are based on Poisson statistics. There is no apparent trend with radius in $N_{\text{sec}}/N_{\text{GSS}}$. The dashed line represents a constant ratio of $N_{\text{sec}}/N_{\text{GSS}} = 0.70$, which is the ratio of stars in the second kinematically cold component to stars in the GSS based on the fits in Table 2. (A color version of this figure is available in the online journal.)

The second cold component, like the GSS, is evenly distributed across the two fields. The bottom panel of Figure 16 shows the ratio of the number of stars in the second kinematically cold component to the number in the GSS ($N_{\text{sec}}/N_{\text{GSS}}$) as a function of R_{proj} . There is no apparent trend in the ratio $N_{\text{sec}}/N_{\text{GSS}}$ with R_{proj} .

A more accurate measure of the velocity dispersions of the two kinematically cold components can be obtained by accounting for the linear trend of line-of-sight velocity with R_{proj} (Figure 16). The line-of-sight velocities of individual stars are shifted to a reference frame defined by the line-of-sight velocity of the GSS as a function of R_{proj} . A maximum-likelihood, triple-Gaussian fit to the data (Section 3.1) is performed. The best-fit values for the observed velocity dispersions of the GSS and second kinematically cold component are $20.1^{+3.8}_{-3.1} \text{ km s}^{-1}$ and $14.1^{+7.1}_{-6.0} \text{ km s}^{-1}$, respectively. These measurements correspond to intrinsic velocity dispersions of 19.4 and 13.0 km s⁻¹, respectively, when the mean radial velocity measurement error in each field is taken into account (Section 3.2.1 and 3.2.2).

There is no evidence in the other GSS core fields of a continuation of the second kinematically cold component seen in fields f207 and H13s (Figure 5). If the linear trend seen in

Figure 16 is extrapolated, the mean velocities expected for the feature in fields IF8 ($R_{\text{proj}} \sim 12$ kpc), IF6 ($R_{\text{proj}} \sim 26$ kpc), and a3 ($R_{\text{proj}} \sim 33$ kpc) are -476 , -344 , and -278 km s $^{-1}$, respectively. The field d1 does have a second kinematically cold component, which is also offset by ~ 100 km s $^{-1}$ from the primary (GSS core) kinematically cold component. It is possible that the second kinematically cold component in fields f207 and H13s is a tidal stream that follows a radial line from fields f207 and H13s to field d1 (and therefore is not present in fields IF6 and a3). However, this interpretation is unlikely based on the metallicities of the stars. Stars in the secondary kinematically cold component in field d1 are uniformly metal rich (Figure 9), and are significantly more metal rich than the second kinematically cold component in fields f207 and H13s ($[\text{Fe}/\text{H}]_{\text{phot}} \sim -0.2$ dex, compared to $[\text{Fe}/\text{H}]_{\text{phot}} \sim -0.5$ dex).

4.3.2. Implications for the Origin of the Second Kinematically Cold Component

The discovery of the second kinematically cold component ($\langle v \rangle = -388$ km s $^{-1}$) in the GSS core field at $R_{\text{proj}} \sim 21$ kpc, field H13s, was originally reported by Kalirai et al. (2006a). These authors considered three possible origins for the secondary kinematically cold component: (1) a feature related to M31's disk; (2) a separate halo tidal stream superimposed on the GSS; and (3) a wrapped around component of the GSS. In this section, we re-evaluate the merits of each of these as the possible physical origin of this feature in light of the newly obtained data in the GSS core field f207 ($R_{\text{proj}} \sim 17$ kpc), the improved statistics in field H13s (Section 2.3), and recent theoretical and observational developments. We also consider a fourth possible origin which was not considered by Kalirai et al. (2006a): a bifurcation in velocity space of the GSS. We will argue that the data in fields f207 and H13s favor a physical connection with the GSS rather than a separate physical origin for the second kinematically cold component in fields f207 and H13s.

M31's stellar disk. The disk of M31 has been observed to large radial distances from the center of M31. An extended smooth stellar distribution has been observed out to ~ 40 kpc in the disk plane, and substructures with disk-like velocities have been observed continuing out in radius to 70 kpc (e.g., Ibata et al. 2005). Furthermore, the mean line-of-sight velocity of the second kinematically cold component is within ~ 50 km s $^{-1}$ of the expected line-of-sight velocity for M31 disk stars in circular orbits at the location of our masks. We therefore consider the possibility that the secondary kinematically cold component in fields f207 and H13s represents a population of smooth disk stars. Kalirai et al. (2006a) show that the surface brightness of the secondary kinematically cold component in field H13s ($R_{\text{disk}} \sim 75$ kpc assuming an unwarped disk of inclination 77° ; Walterbos & Kennicutt 1988) is more than 25 times brighter than would be expected based on an extrapolation of M31's smooth exponential disk, effectively ruling out the smooth disk of M31 as the origin of the substructure (Figure 11 of Kalirai et al. 2006a). The surface brightness of the secondary peak in field f207 ($R_{\text{disk}} \sim 59$ kpc) is equivalent to that of the secondary kinematically cold component in field H13s, and is approximately 10 times brighter than the extrapolation of M31's smooth disk. The surface brightness of M31's disk decreases exponentially with increasing projected radius while the surface brightness of the GSS is relatively constant over the radial range covered by the f207 and H13s spectroscopic masks. The f207 and H13s masks combined cover a distance of ~ 30 kpc in the plane of the disk, which is more than five scale lengths

(Walterbos & Kennicutt 1988). Thus, the ratio of the number of stars in the second kinematically cold component to the number of stars in the GSS would be expected to decrease substantially with increasing projected radius if the second kinematically cold component was part of M31's disk. This is contrary to what is observed (Figure 16(b)). These arguments rule out the possibility that the origin of the second kinematically cold component in these fields is M31's extended stellar disk.

Unrelated substructure. Cosmological simulations of halo formation (e.g., Bullock et al. 2001; Bullock & Johnston 2005) as well as observational studies of the MW (e.g., Majewski et al. 1996; Newberg et al. 2002; Belokurov et al. 2006) suggest that multiple tidal streams may exist along a single line of sight. The secondary kinematical peaks in fields f207 and H13s could represent a tidal stream in M31's halo unrelated to the progenitor of the GSS. Also, as mentioned above, substructures rotating with disk-like velocities have been observed out to $R_{\text{disk}} \sim 70$ kpc from M31's center (Ibata et al. 2005). The arguments against a smooth disk origin for the second kinematically cold component in fields f207 and H13s do not preclude the possibility that it is substructure related to the disk. However, for either halo or disk-related substructures we would not expect a direct correlation between the line-of-sight velocity of the GSS and the second stream. Since the two kinematical substructures are superimposed on each other in projected location, only a limited set of orbits would produce the observed correspondence in mean line-of-sight velocity with R_{proj} (Figure 16).

Wrapped around GSS component. Kalirai et al. (2006a) discussed the possibility that the second kinematically cold component in field H13s represented part of a later orbital phase of the GSS's progenitor that has wrapped around M31 and back to the location of this field. A few of the many dynamical models of the merging event (Ibata et al. 2004; Font et al. 2006b; Fardal et al. 2006) yield a progenitor orbit that passes through the region of fields H13s and f207 twice (e.g., Figure 8 of Fardal et al. 2006). This origin was disfavored by Kalirai et al. (2006a) because the stars are predicted to be moving in very different phases during the two passages. This results in a predicted line-of-sight velocity offset between the stellar streams produced by each phase that is inconsistent with the observed velocity offset between the GSS core and the second kinematically cold component in fields H13s and f207 (Table 2). The most recent dynamical models of the collision suggest that the NE and W shelves observed in M31 are the forward continuation of the GSS (Fardal et al. 2007). This model is strongly supported by recent spectroscopic and *HST*/ACS observations of the shelves in M31 (Brown et al. 2006b; Gilbert et al. 2007; Richardson et al. 2008). Although in this model the progenitor's orbit does not pass through the region of fields H13s and f207 a second time, the shelves themselves could be responsible for the second kinematically cold component. The extent of the W shelf is obscured by the southwest region of M31's disk. It could produce a velocity peak with a ~ 100 km s $^{-1}$ separation from the GSS velocity in fields f207 and H13s, although the second velocity peak characteristic of shelf features (one peak corresponds to outflowing and one to infalling debris) would have to be depopulated to match the data (M. Fardal 2009, private communication). The W shelf would have to span almost 180° in position angle to be the source of the kinematically cold component in fields f207 and H13s. Future spectroscopic observations will test whether the velocity of W shelf stars are consistent with the second kinematically cold component in the GSS core fields f207 and H13s.

Bifurcation of the GSS. The fact that the velocities of stars in the GSS and second kinematically cold component follow the same linear trend with projected distance from M31's center, and that the ratio of stars in the two kinematical features is approximately constant with R_{proj} , suggests the possibility of a direct physical association between the two features. The formation of velocity caustics during the tidal disruption of an infalling satellite can lead to bifurcations in the observed line-of-sight velocities of a tidal feature (e.g., Merrifield & Kuijken 1998; Fardal et al. 2007). Thus, the second kinematically cold component in fields f207 and H13s could be part of the GSS itself (i.e., stars, like those in the GSS, that are approaching their first pericentric passage since the progenitor's initial encounter with M31). However, it may be difficult to produce a bifurcation with a separation of line-of-sight velocities as large as $\sim 100 \text{ km s}^{-1}$, as seen in fields f207 and H13s. Further modeling of the collision will be needed to determine whether velocity caustics of this magnitude can be produced in this region of the GSS.

If the second kinematically cold component is related to the GSS, either as a bifurcation of the GSS itself or the forward continuation of the GSS, the metallicity distributions of stars in the GSS and the second kinematically cold components should be similar. Figure 17 displays the cumulative $[\text{Fe}/\text{H}]$ distributions of stars in the GSS core (solid curve) and the second kinematically cold component in field f207 (dotted curve) and H13s (dashed curve). Although the second kinematically cold components have a small number of stars more metal rich than any associated with the GSS core, the $[\text{Fe}/\text{H}]$ distributions are not inconsistent. A two-sided KS test returns probabilities of 15% and 22% that the $[\text{Fe}/\text{H}]$ distributions of the second kinematically cold components in f207 and H13s, respectively, were drawn from the same parent distribution the GSS core was drawn from.

In conclusion, the data rule out M31's extended, smooth disk as the origin of the second kinematically cold components in fields f207 and H13s. While substructure in M31 unrelated to the GSS cannot be ruled out by the data, the possibility of a connection with the GSS, either via a bifurcation in velocity space of the GSS itself or a wrapped around component of the GSS, is an intriguing possibility. The common trend of mean line-of-sight velocity with projected radius of the GSS and the second kinematically cold component combined with the similar metallicity distributions of their stars would be a remarkable coincidence if the substructures are unrelated, but are naturally explained if the second component is also part of the GSS.

4.4. Summary of Debris Properties

We have detected the metal-rich GSS core in fields f207, H13s, a3, and d1. Although field d1 is located on the dSph And I, we are still able to isolate M31 field stars from dSph members, and to identify kinematically cold features in the M31 field star population. The kinematical and chemical properties of the -380 km s^{-1} kinematically cold component in field d1 are entirely consistent with its identification as GSS core debris. The kinematical and chemical properties of the relatively metal-poor -300 km s^{-1} cold component in field a13 are consistent with what is expected for the GSS envelope based on empirical evidence from photometric studies and theoretical modeling of the collision that produced the GSS. The second kinematically cold component in fields f207 and H13s has a similar $[\text{Fe}/\text{H}]_{\text{phot}}$ distribution to the GSS core, follows the same parallel trend in the line-of-sight velocity versus R_{proj} plane, and may also be associated with the GSS.

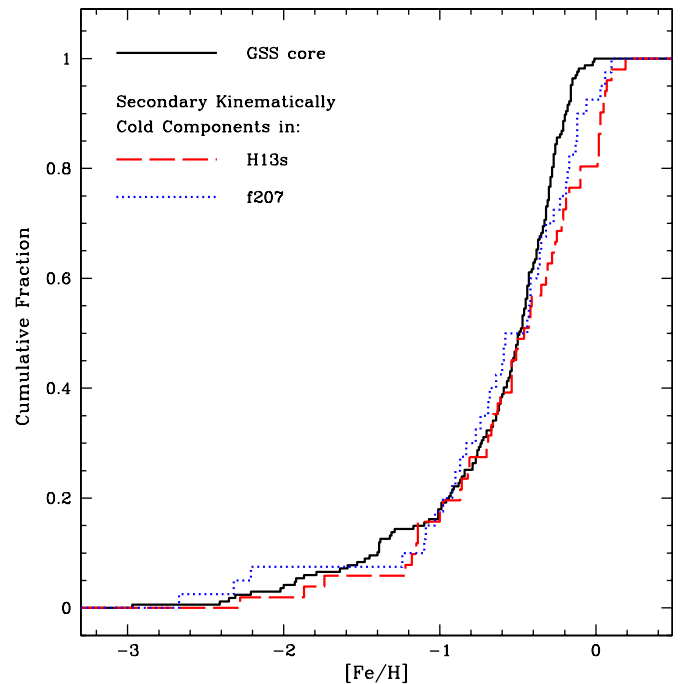


Figure 17. Cumulative $[\text{Fe}/\text{H}]$ distribution of stars within $\pm 2\sigma_v$ of the mean velocity of the second kinematically cold components in fields H13s (dashed curve) and f207 (dotted curve). The cumulative $[\text{Fe}/\text{H}]$ distribution of the GSS core (based on the primary kinematically cold component in fields f207 and H13s and the kinematically cold component in the a3 field) is shown for comparison (solid curve; Figure 15). The second kinematically cold component in fields f207 and H13s is as metal rich as the GSS core.

(A color version of this figure is available in the online journal.)

The kinematically cold features observed in field m4 confirm the identification by Chapman et al. (2008) of two kinematically cold components (one metal rich and one metal poor) in the photometric stream C minor-axis feature. The origin of the more metal-rich kinematical feature is not known. The more metal-poor feature may be physically associated with the extended stellar cluster EC4, although it is unlikely to be a tidal stream originating from EC4 itself.

5. DISCUSSION

Mixed stellar populations appear to be an inevitable feature of the M31 halo. Many fields in the SPLASH spectroscopic survey show conclusive (Table 2) or suggestive evidence of multiple components within the M31 field halo population itself. In addition, fields that target M31's dSph galaxies contain dSph member RGB stars as well as M31 field halo RGB stars (Section 2.5.1). All fields contain a non-negligible fraction of foreground MW dwarf star and background galaxy contaminants. It is useful to keep the notion of mixed populations in mind while interpreting star-count maps. In this section, we discuss several implications of the mixed stellar populations discovered in M31 for our understanding of the structure and formation of M31's stellar halo.

Although many of the spectroscopic fields presented here specifically targeted tidal debris (f207, H13s, IF8, IF6, IF2, and IF1), others were chosen for DEIMOS spectroscopy without prior knowledge of the existence of tidal debris at their locations (fields a3, m4, a13, d1, and d3). Nevertheless, the majority of stars are found to be located in distinct clumps in $v_{\text{los}}-[\text{Fe}/\text{H}]$ space. Moreover, many fields contain multiple distinct groupings of stars. If the outer regions of M31's stellar halo were

well mixed, stars from numerous progenitors would be present along any given line of sight, leading to stars that are broadly distributed in $v_{\text{los}} - [\text{Fe}/\text{H}]$ space. Tremaine (1993) calculates that the filling factor of tidal streams, f_3 , in a MW type halo built from the disruption of accreted satellites should be small ($f_3 < 1$) at radii of ~ 50 kpc (a given volume of the halo is dominated by stars from one to a few progenitors), while well-mixed systems have $f_3 \gg 1$ (stars from numerous progenitors are present in any given volume). The fields presented above indicate that a large fraction of the stars in M31's halo at these radii belong to groups of stars that can be identified as individual tidal streams. As discussed by Majewski et al. (1996), this may prove problematic for drawing conclusions about the global properties of M31's stellar halo based on limited lines of sight. Further studies in other quadrants of M31's stellar halo will be able to determine whether the dominance of distinct groups of stars in these fields is a ubiquitous property of M31's stellar halo, or is mainly a result of the large amounts of debris deposited by the progenitor of the GSS.

The number of distinct kinematical components present in many of the M31 stellar halo fields necessitates careful interpretation of the properties of specific regions and tidal features measured from photometric studies of M31's halo population. Even if the percentage of the population associated with a tidal feature is known via separate kinematical studies, photometric studies cannot identify which stars are more likely associated with a specific feature rather than the underlying halo of M31. Metallicity distributions of features derived purely from photometric observations will therefore be a sum of the metallicity distribution of the feature itself, M31's halo, and any other structural (sub)component or feature present in the analyzed region. The minor-axis arc stream C provides a prime example of the complexity that can be present in a small region of M31's halo (Sections 3.3.2 and 4.2). Furthermore, the extensive presence of mixed stellar populations in M31's stellar halo, seen in both kinematical (this work; also see Ibata et al. 2004, 2005; Guhathakurta et al. 2006; Kalirai et al. 2006a; Gilbert et al. 2007) and photometric surveys (e.g., Ibata et al. 2001, 2007; Ferguson et al. 2002), makes identifying control M31 halo fields for photometric studies problematic. Kinematical data provide the ability to identify which stars are likely to be associated with an individual (sub)component or tidal feature; this provides a distinct advantage over photometric studies in the analysis of the properties of individual stellar (sub)components in M31.

A variety of spatial structures are expected in stellar halos that form via the merging of smaller stellar systems, including still-bound satellites, spatially coherent tidal streams, larger, more diffuse clouds of debris, and a spatially smooth, well-mixed stellar population. The data presented here display evidence of a correlation between kinematical and spatial structure. The M31 dSphs And I and And III are very kinematically cold, with low velocity dispersions ($\sigma_v < 10 \text{ km s}^{-1}$; J. S. Kalirai et al. 2009, in preparation). As surviving dSphs, these objects are also spatially compact. Many of the kinematically cold components observed in our spectroscopic fields (i.e., m4 and the GSS core in f207, H13s, and a3) can be identified with spatially sharp debris streams. The kinematically cold component in field a13 is warmer (has a broader velocity distribution) than the kinematically cold components in the GSS core fields. It is identified with the GSS envelope debris, which appears to be a broad cloud of debris rather than a sharp tidal stream. Our fields also contain a kinematically hot population with a very

broad spread in velocities; this population represents M31's underlying, well-mixed stellar halo.

Metal-rich tidal streams are expected to dominate over metal-poor streams in surface-brightness-limited stellar halo surveys, due primarily to the fact that more massive dwarf galaxies are both more metal rich (e.g., Larson 1974; Mateo 1998; Dekel & Woo 2003) and tend to produce higher-surface-brightness stellar streams (Johnston et al. 2008; Gilbert et al. 2009). We currently have $\lesssim 50$ M31 RGB stars per spectroscopic field in M31's outer halo ($R_{\text{proj}} \gtrsim 35$ kpc). A high-surface-brightness tidal stream is easily identifiable in our survey in the form of a kinematically cold component seen against the dynamically hot underlying stellar halo. Tidal streams with a low level of contrast against M31's spheroid are difficult to detect due to our limited sample size. We therefore expect the kinematically cold components identified in our M31 fields to be more prominent among the metal-rich than metal-poor stars. In general, this is indeed the case, as shown in panels (b)–(d) of Figures 6–12.

6. CONCLUSIONS

In this study, we presented data for five new spectroscopic fields observed during the course of our SPLASH Keck/DEIMOS survey. We also presented improved reductions and sample selection for two of our previously published fields (a3 and H13s; Guhathakurta et al. 2006; Kalirai et al. 2006a) and compiled all previously published line-of-sight velocity data in fields on the GSS. This combined spectroscopic data set was used to analyze the kinematical and chemical properties of the GSS.

The main findings of this paper are as follows.

1. We have detected and measured the mean velocity and velocity dispersion of the GSS in two new GSS core fields: (1) field f207, at $R_{\text{proj}} \sim 17$ kpc (the innermost kinematic measurement of the GSS to date), and (2) field d1 at $R_{\text{proj}} \sim 45$ kpc (the first detection of GSS core debris in a field well removed from the sharp eastern edge of the GSS).
2. We have made the first kinematical detection and measurement of the mean velocity and velocity dispersion of the GSS envelope in field a13, at $R_{\text{proj}} \sim 58$ kpc.
3. We have presented updated measurements of the mean velocity and velocity dispersion of two previously published GSS core fields at $R_{\text{proj}} \sim 21$ and 33 kpc using improved reductions of the spectroscopic data.
4. The full data set of new and previously published line-of-sight velocities in GSS fields was used to measure the dependence of the mean line-of-sight velocity and velocity dispersion of the GSS as a function of projected distance from M31's center.
5. We utilized the line-of-sight velocities in the SPLASH survey fields to identify stars belonging to the GSS. The metallicity distribution of stars in the GSS is found to be similar in all GSS core fields, while the mean metallicity of stars in the GSS envelope is significantly more metal poor (~ 0.7 dex) than the GSS core.
6. We have confirmed the presence of two kinematically cold components, one relatively metal rich and the other relatively metal poor, in the stream C region at $R_{\text{proj}} \sim 60$ kpc on the minor axis. A comparison of the

mean velocities of the kinematic features in field m4 with a previously published field 1° to the southwest indicates that the observed line-of-sight velocities of the two features diverge as one moves from the southwest to the NE along stream C.

7. In our innermost GSS core field ($R_{\text{proj}} \sim 17$ kpc), we discovered the continuation of a second kinematically cold component previously seen in a GSS core field at $R_{\text{proj}} \sim 21$ kpc. The kinematical and chemical properties and number of stars in this second kinematically cold component are tightly correlated with those of the GSS over $\Delta R_{\text{proj}} = 7$ kpc, indicating that a direct physical association between it and the GSS is likely. The second kinematically cold component in these two fields may represent a bifurcation in the line-of-sight velocities of stars in the GSS.

Previous spectroscopic data on the GSS have been presented piecemeal in several different papers (Ibata et al. 2004; Guhathakurta et al. 2006; Kalirai et al. 2006a). The compilation of previously published data and the new GSS fields presented here provide a clear portrait of the kinematical and chemical properties of the GSS along a significant portion of both its length and breadth. This data provide ingredients for improved detailed modeling of the collision of the progenitor of the GSS with M31, and will provide additional observational constraints on M31's potential as well as the properties and orbit of the progenitor of the GSS.

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